



# 3 $\mu\text{m}$ 帯同時分光で探る 原始惑星系円盤、及び系外惑星大気

寺田 宏

(国立天文台ハワイ観測所)

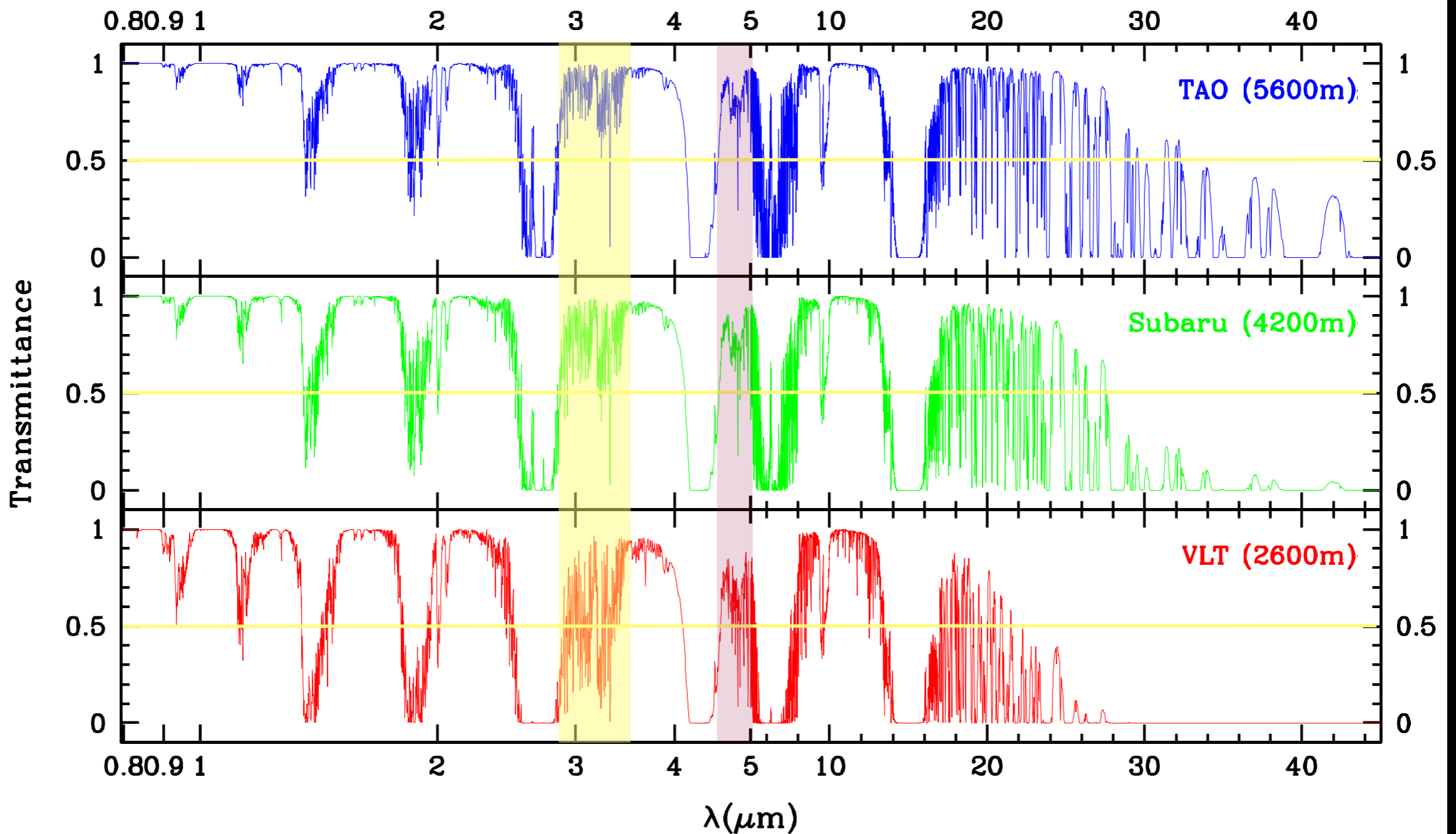
- (I) *Organic Molecules and Ices in Protoplanetary Disks*
- (II) *Exo-Planetary Atmosphere in Transiting Systems*



# Advantage of MIMIZUKU

*"Great Site" & "Simultaneity"*

*Atmospheric Calibration for Less Transmittance IR-Region*



(I) Organic Molecules  
and Ices in  
Protoplanetary Disks



# Observational View of Protoplanetary Disks

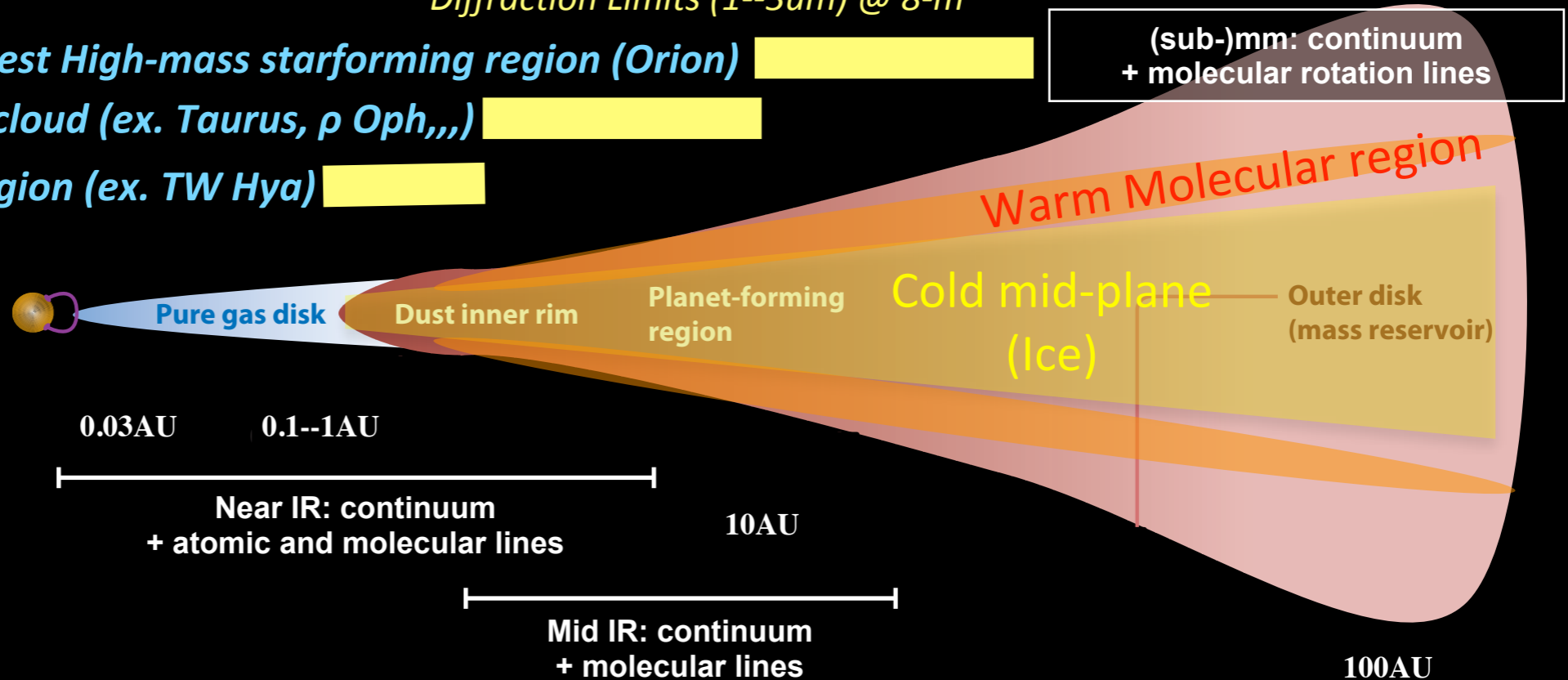
Diffraction Limits (1--3um) @ 8-m

Nearest High-mass starforming region (Orion)

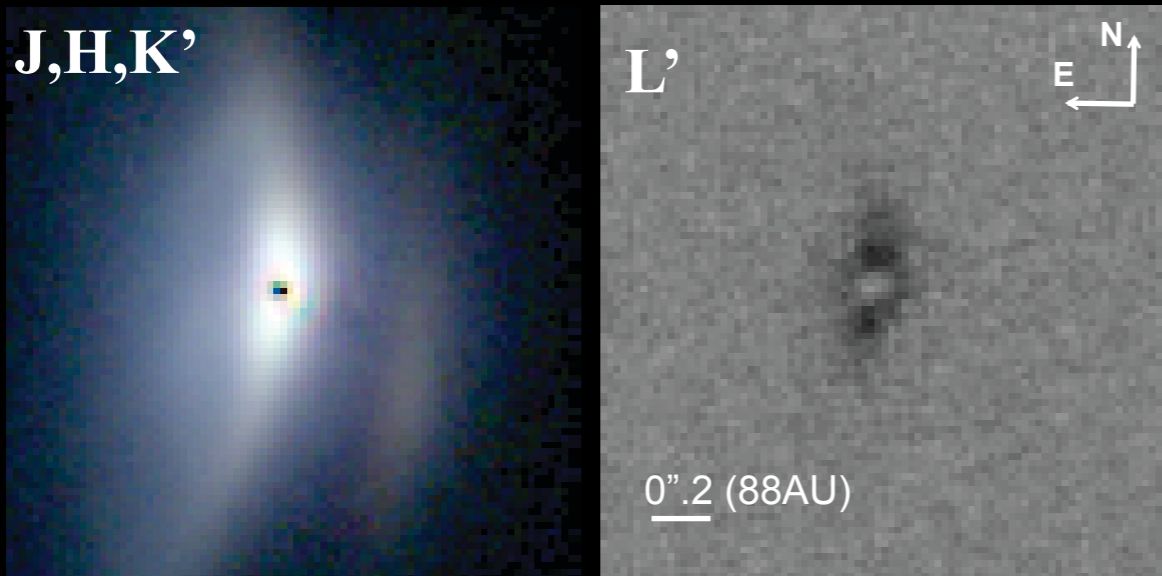
Nearby molecular cloud (ex. Taurus,  $\rho$  Oph,,,) )

Nearest starforming region (ex. TW Hya)

(sub-)mm: continuum  
+ molecular rotation lines



IRCS + AO188 [NGS] @ Orion



Ground-based AO works well  
in NIR region.

=> Spectro-Astrometry w/ AO  
is a hope to probe the inner  
region..

Terada+ 2012



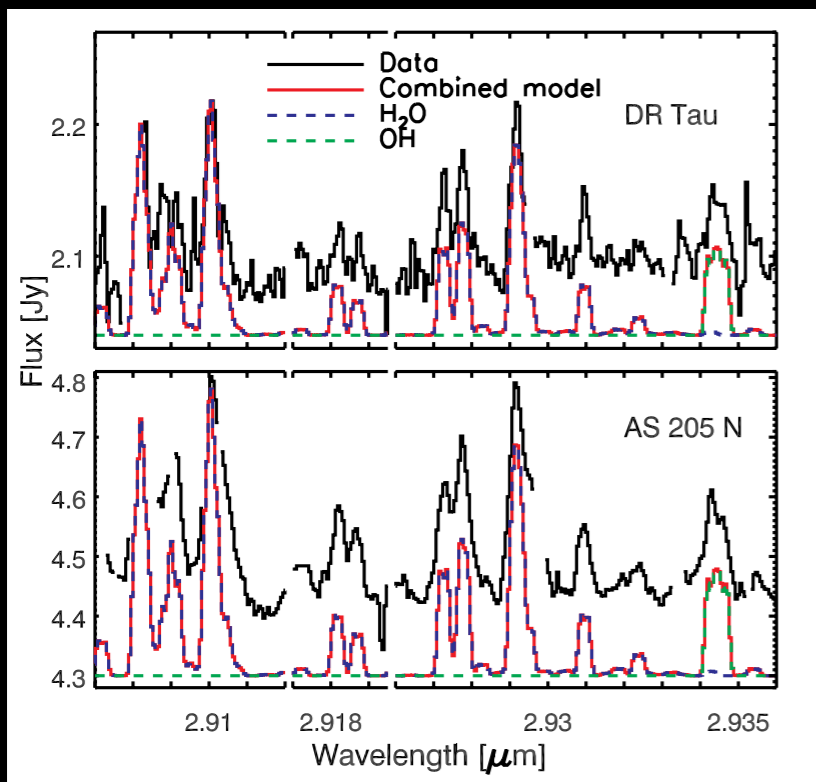
# (a) Organic Molecules in Protoplanetary Disks

## Previous Detections

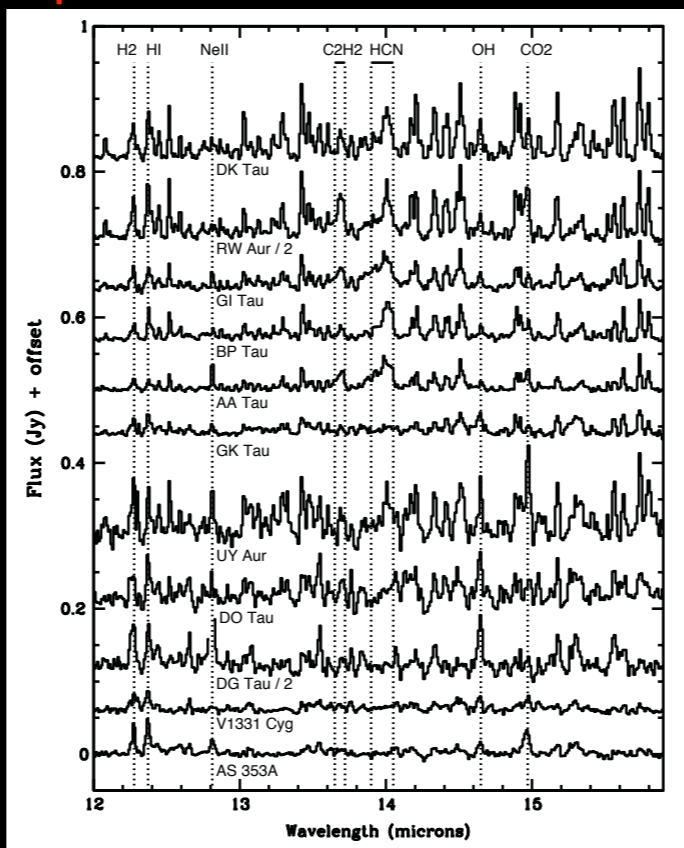
Spitzer C<sub>2</sub>H<sub>2</sub> HCN OH CO<sub>2</sub>

Herschel

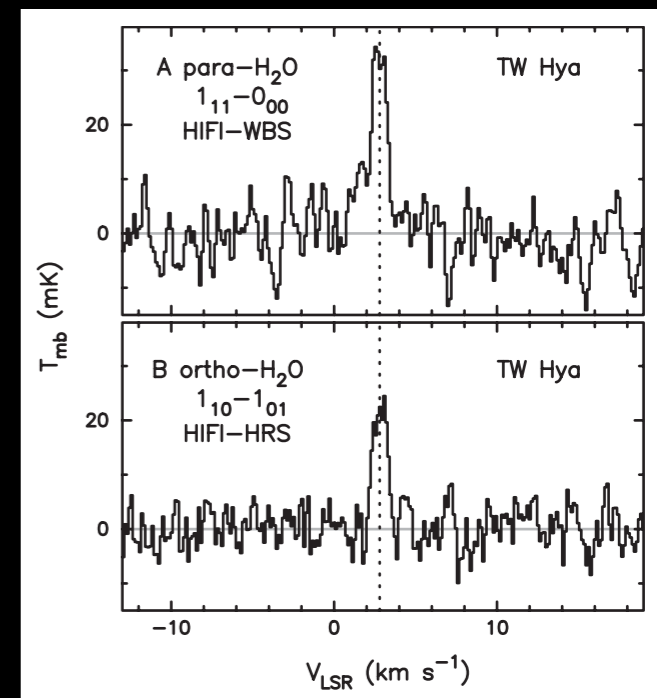
Keck NIRCPEC



Salyk+ 2008

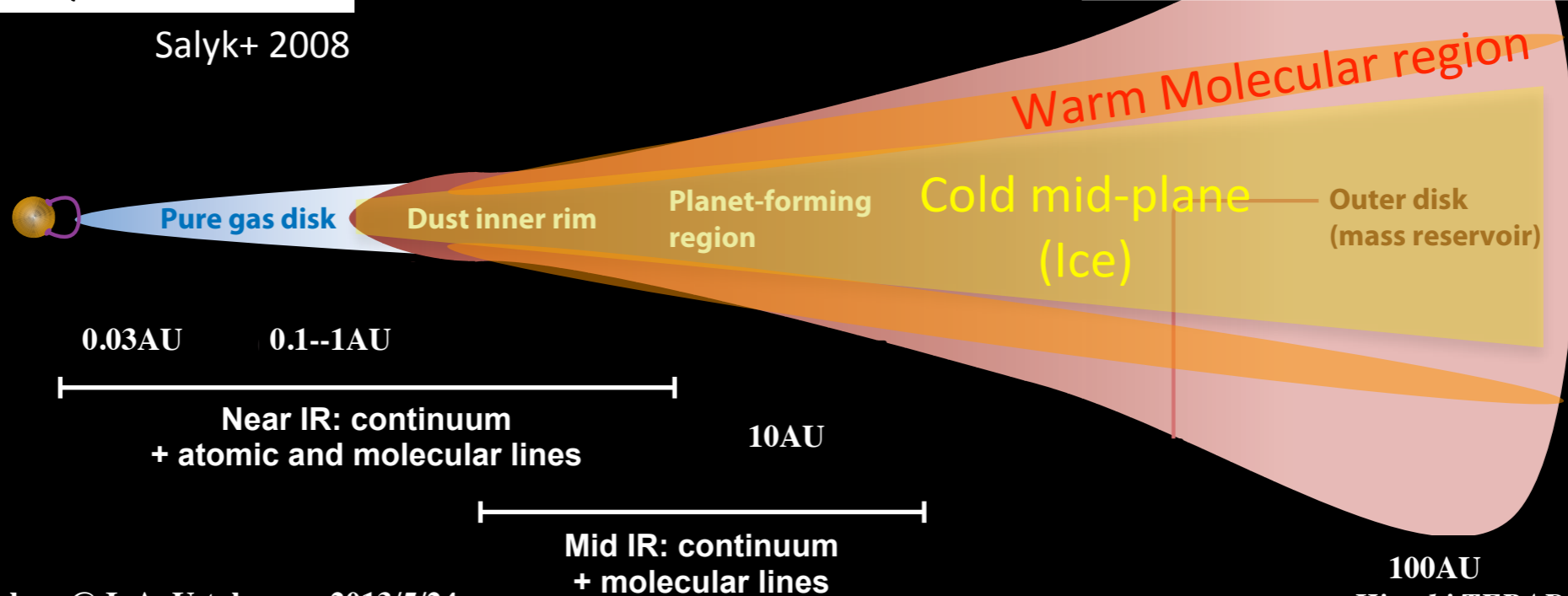


Carr+ 2011



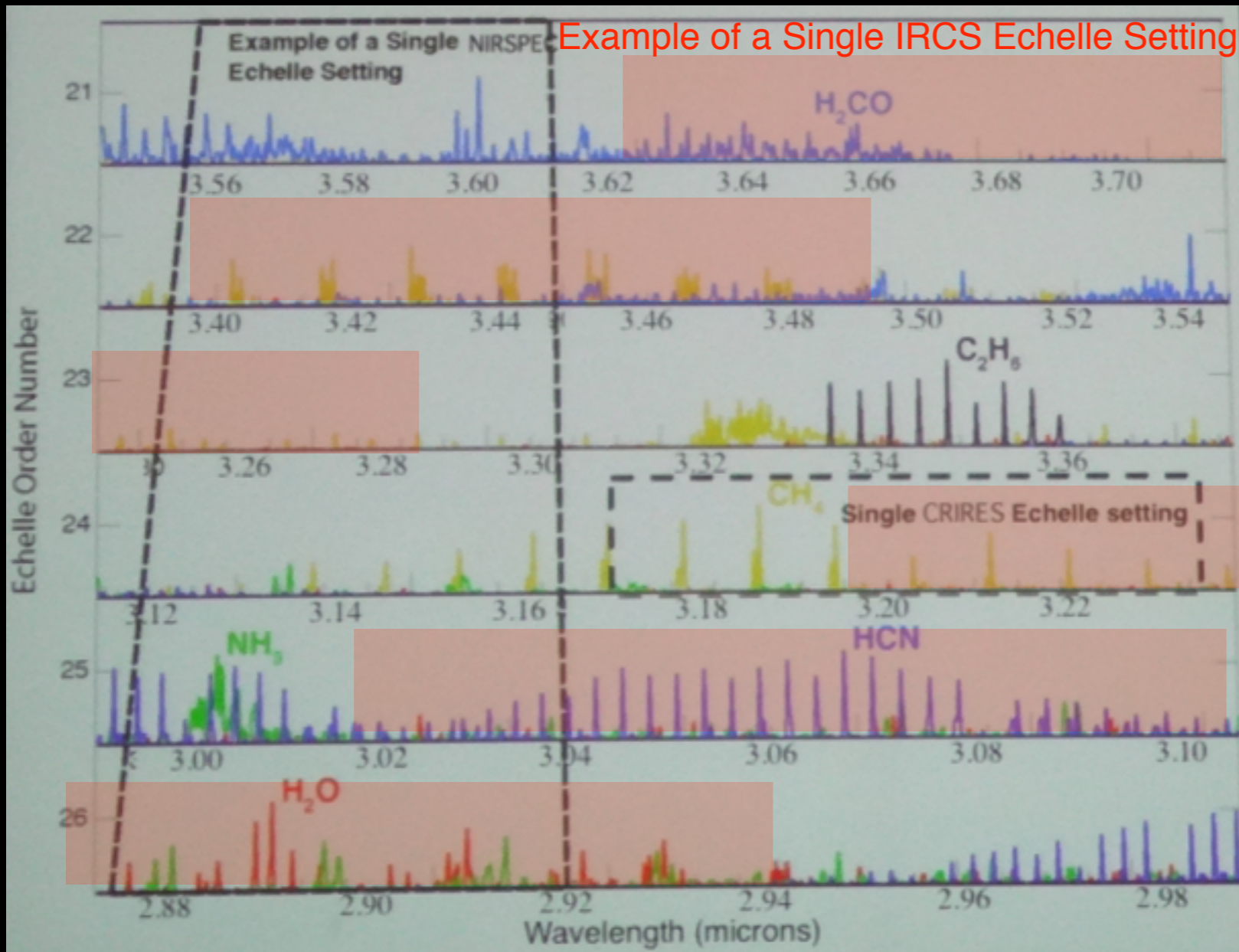
Hogerheijde+ 2011

(sub-)mm: continuum + molecular rotation lines





# (a) Organic Molecule Signatures *Near IR (~3um) Wavelengths*



*Suitable for tracing  
warm molecular volatile  
gases.*

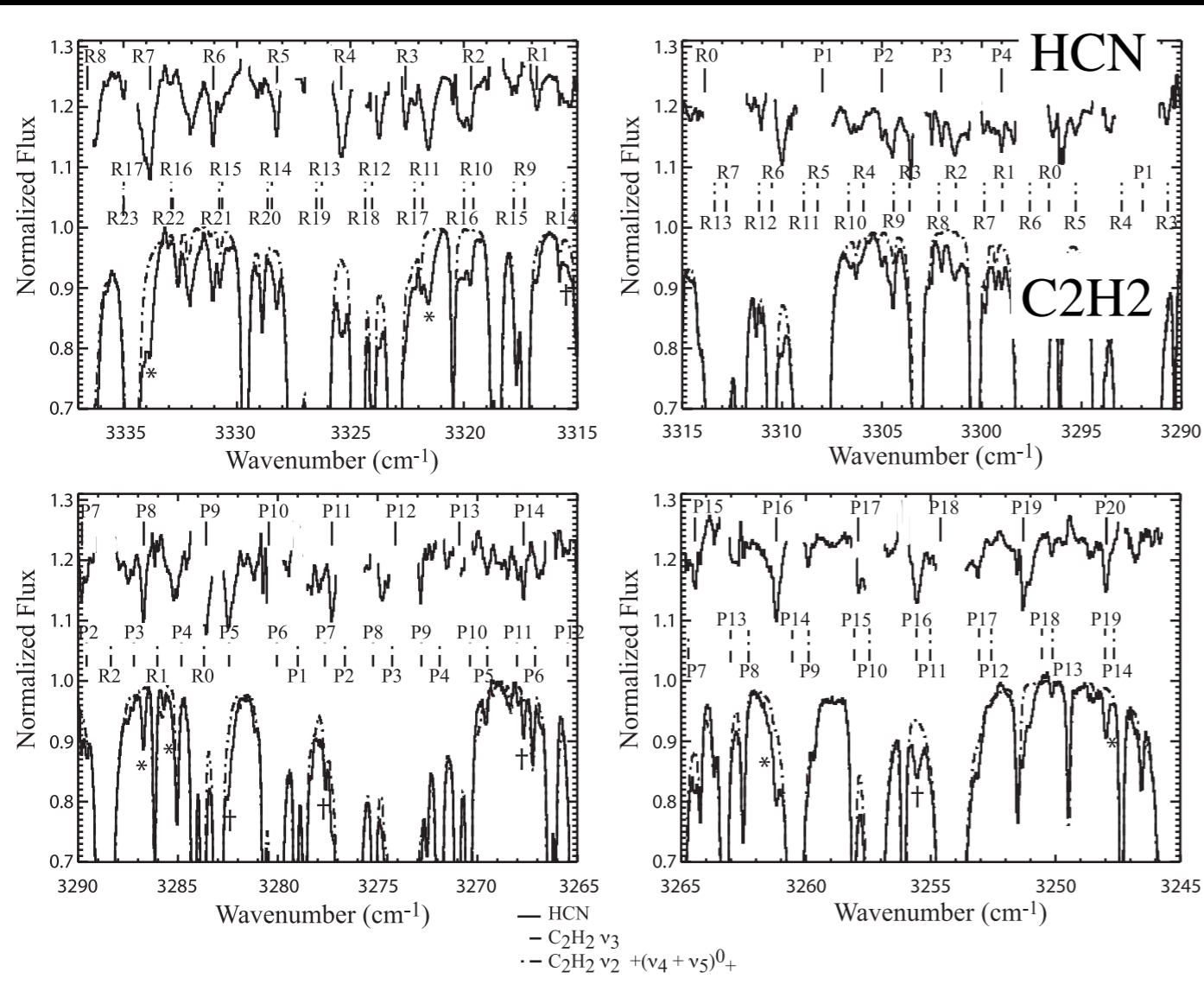
*Sensitive for temperature  
(100--1000K)  
=> <5AU molecule*

*Limitation:  
Atmospheric correction*

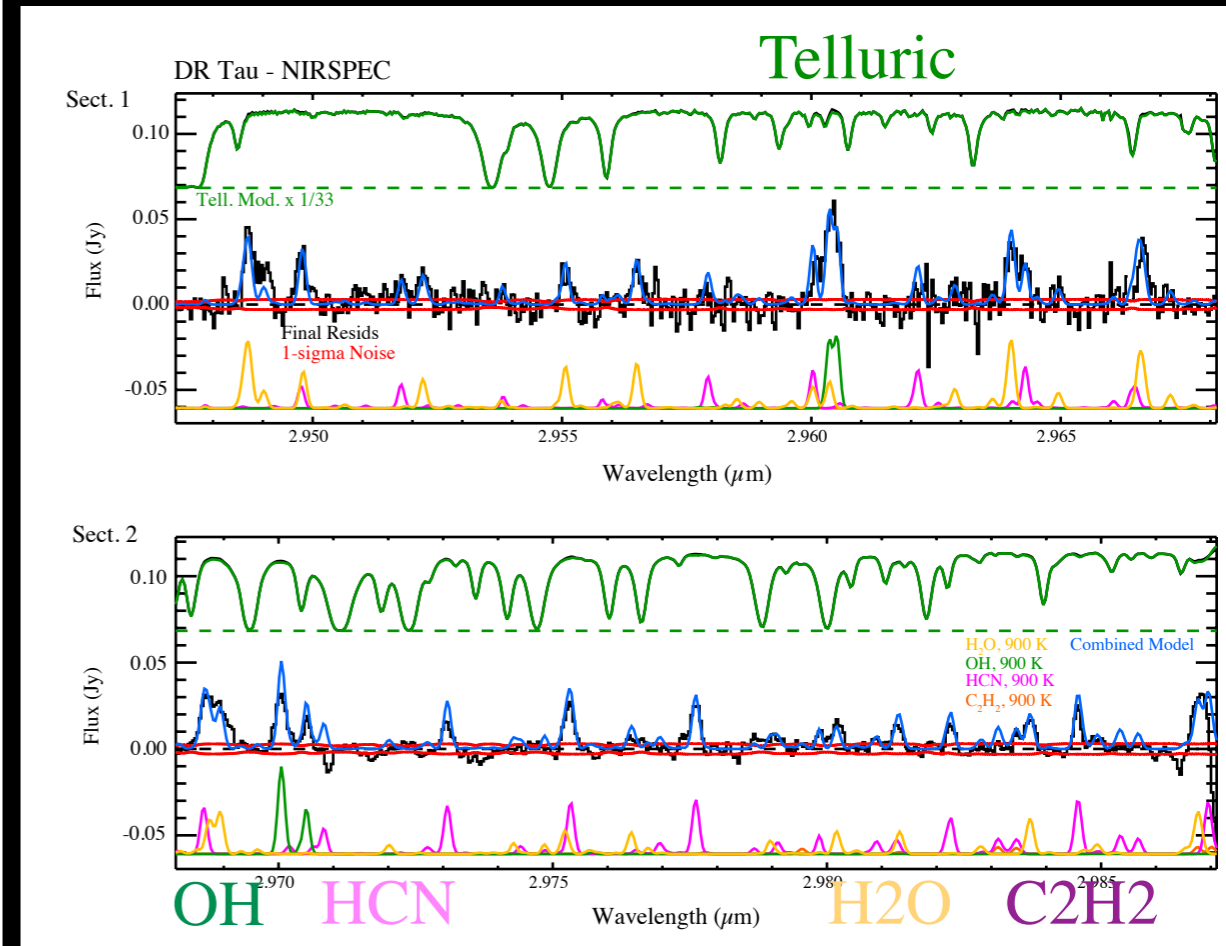
Background figure from Mandell 2012



# (a) Organic Molecules in Protoplanetary Disks: *Difficulties in Detection*



Gibb+ 2007



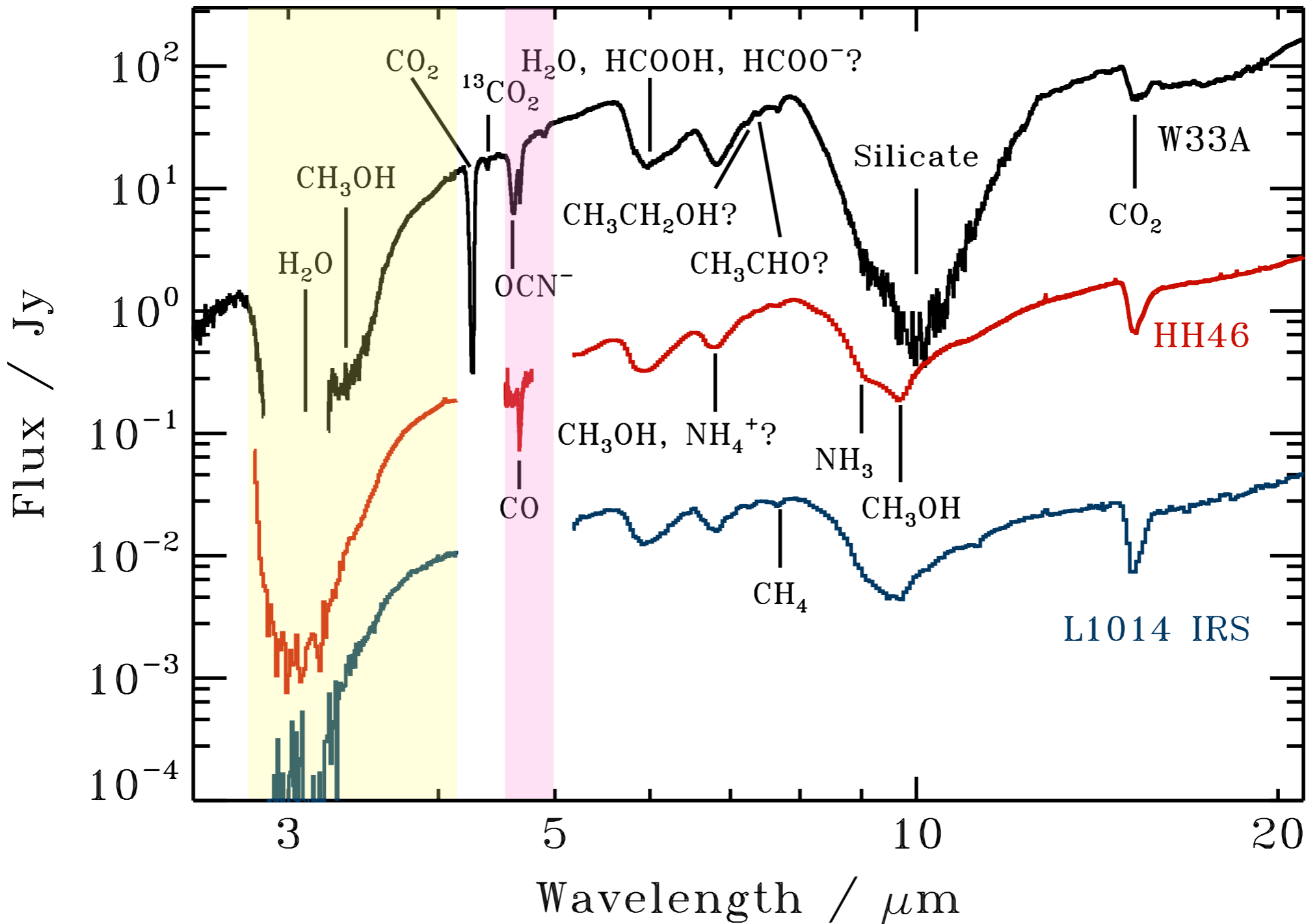
Mandell+ 2012

=> *Higher resolution is definitely better*



# (b) Ices Signatures Towards Protostars

Oberg+ 2011



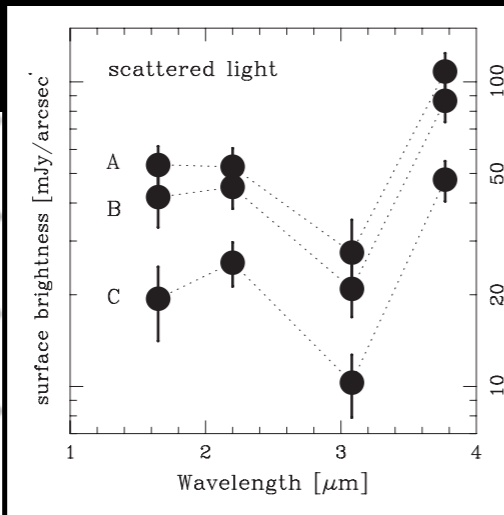
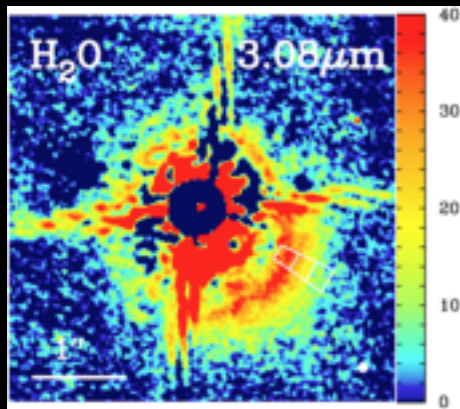




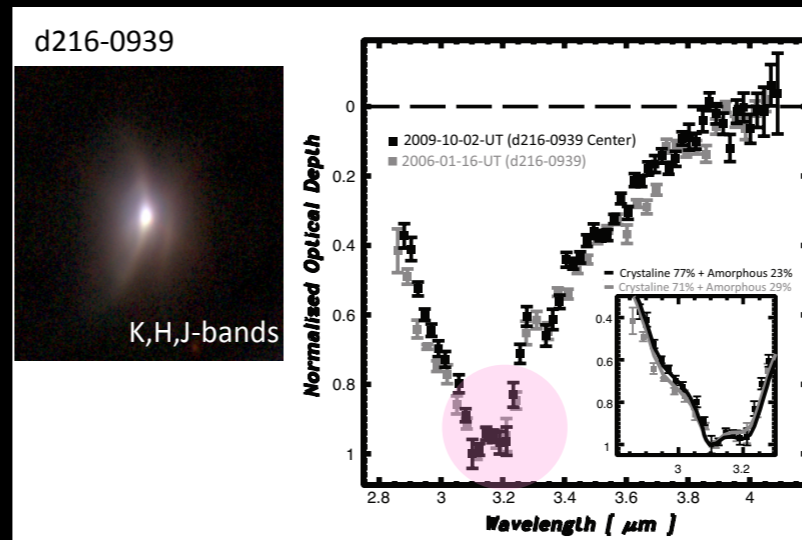
# (b) Ices in Protoplanetary Disks

## Previous Detections

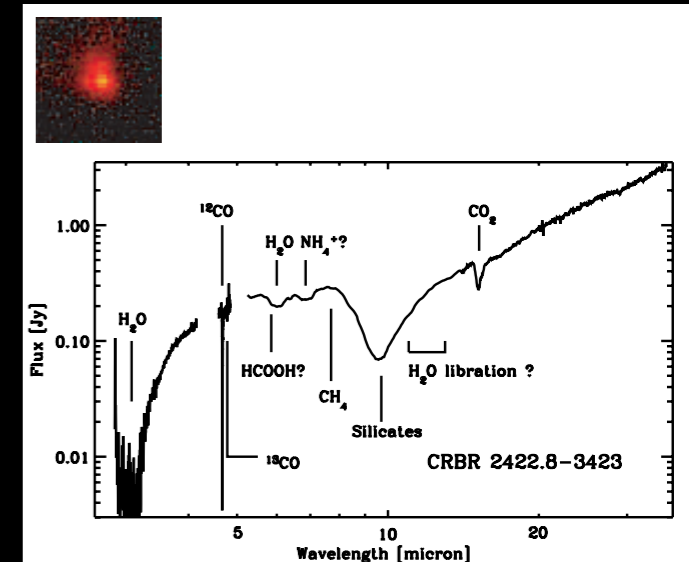
Honda+ 2009



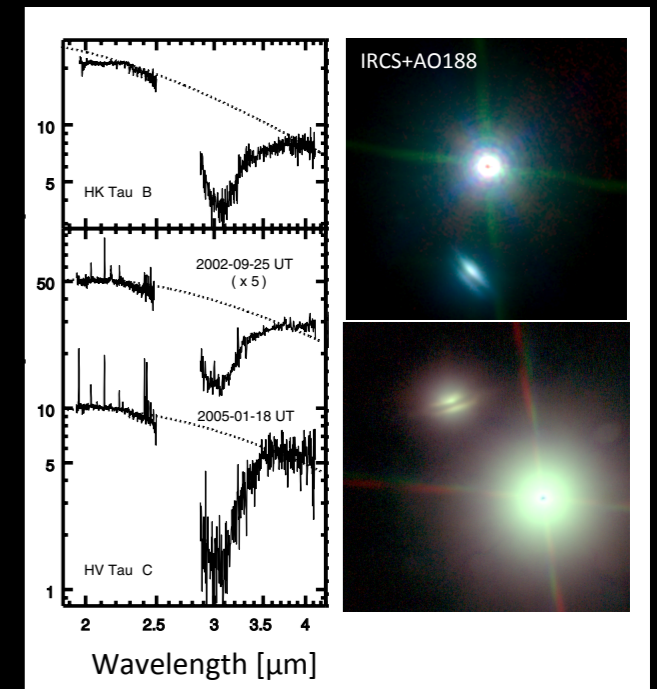
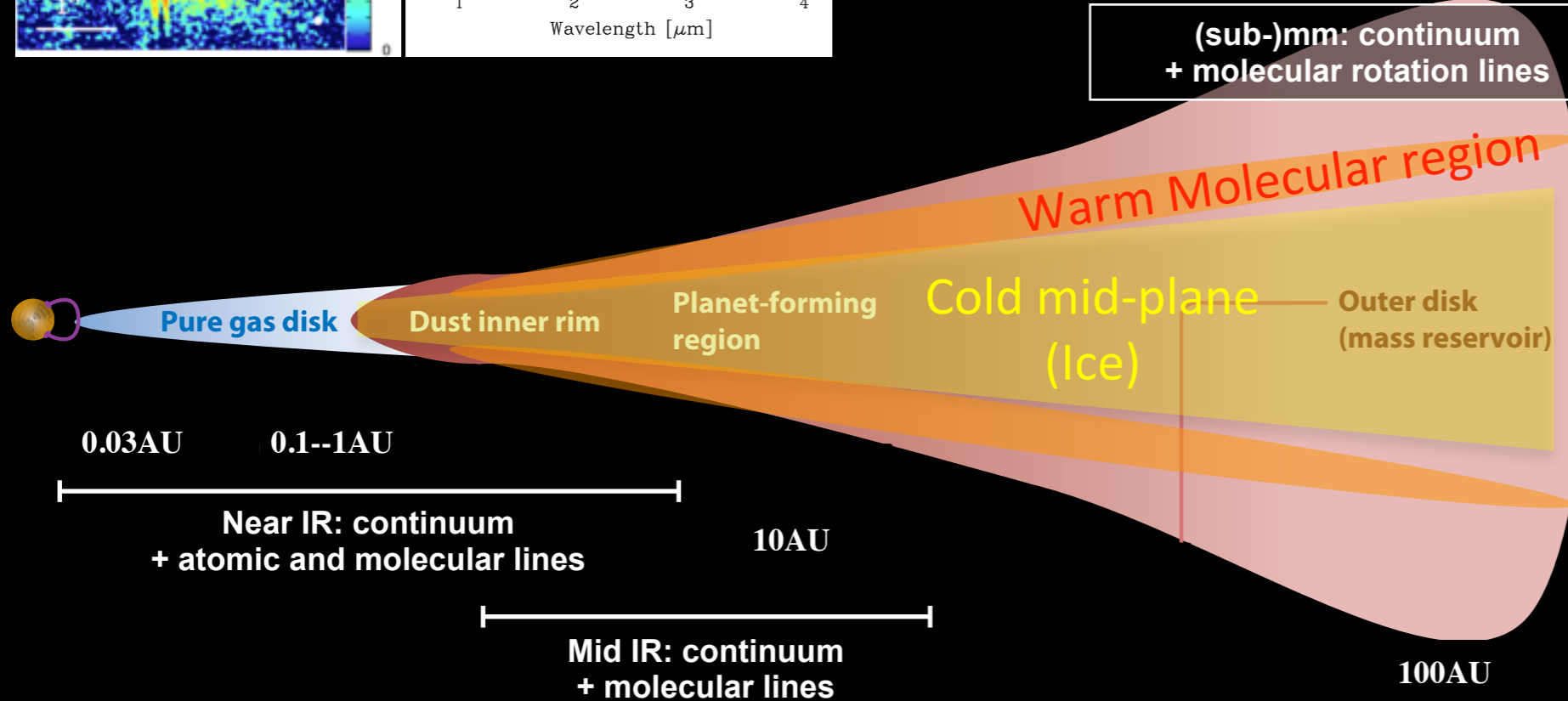
Terada+ 2012



Pontoppidan+ 2005



(sub)-mm: continuum + molecular rotation lines

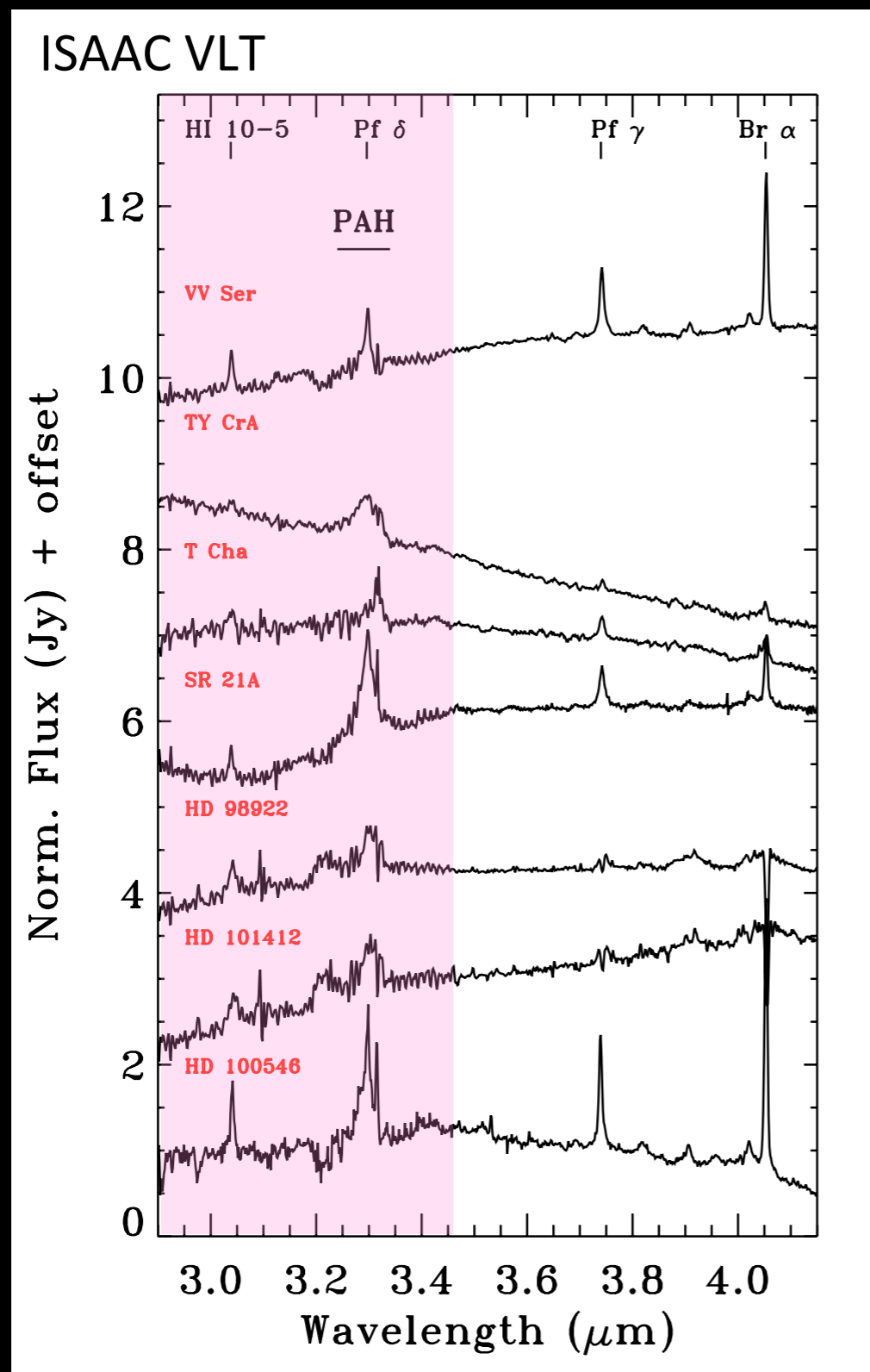


Terada+ 2007



# (b) Ices Signatures

## *Difficulties in Detection*



@ VLT site

*Worse quality in  $< 3.5\mu\text{m}$  region  
even for such bright objects ( $L\sim 5$ )*

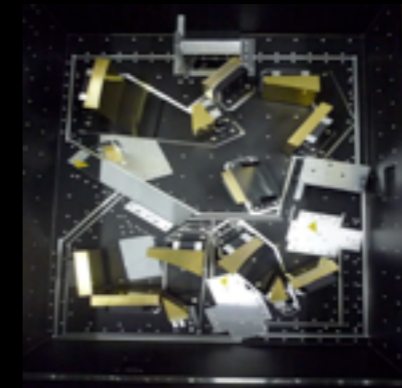
*Very difficult for its profile  
discussion...*

Geers+ 2007

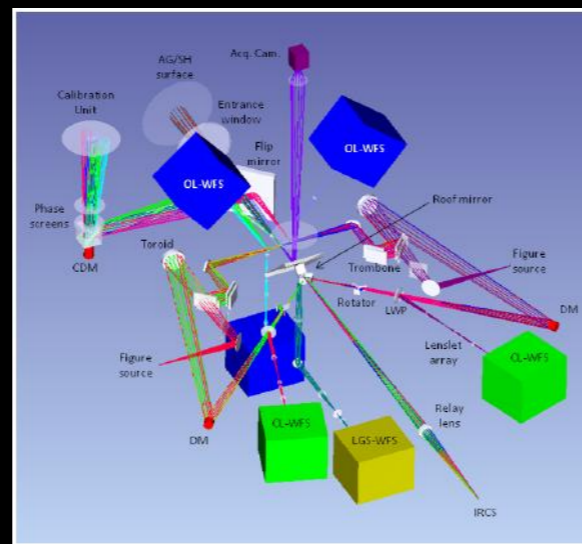


# Facilities for 3 $\mu$ m Simultaneous Spectroscopy *for BEST cancelation of telluric absorption*

## ◆ *MIMIZUKU @ Subaru, TAO*



## ◆ *RAVEN (+IRCS) @ Subaru [2014 Spring]*



“2” Science Path  
MOAO system

UVic, HIA, Subaru, Tohoku-U



# Complementarity for 3 $\mu$ m Simultaneous Spectroscopy

## ◆ *MIMIZUKU @ Subaru, TAO*

- *Patrol area: ~6', ~25'*
- *Fully optimized for thermal-IR bands*
- *Low resolution (R=180) only*

*Ice*

## ◆ *RAVEN (+IRCS) @ Subaru [2014 Spring]*

- *Diffraction limited image*
- *High resolution (R=20,000) is also available*
- *Slit PAs can be independent.*
- *Patrol area: 3'.5*
- *Not optimized for thermal-IR bands*
- *Need "2" NGS guide stars*

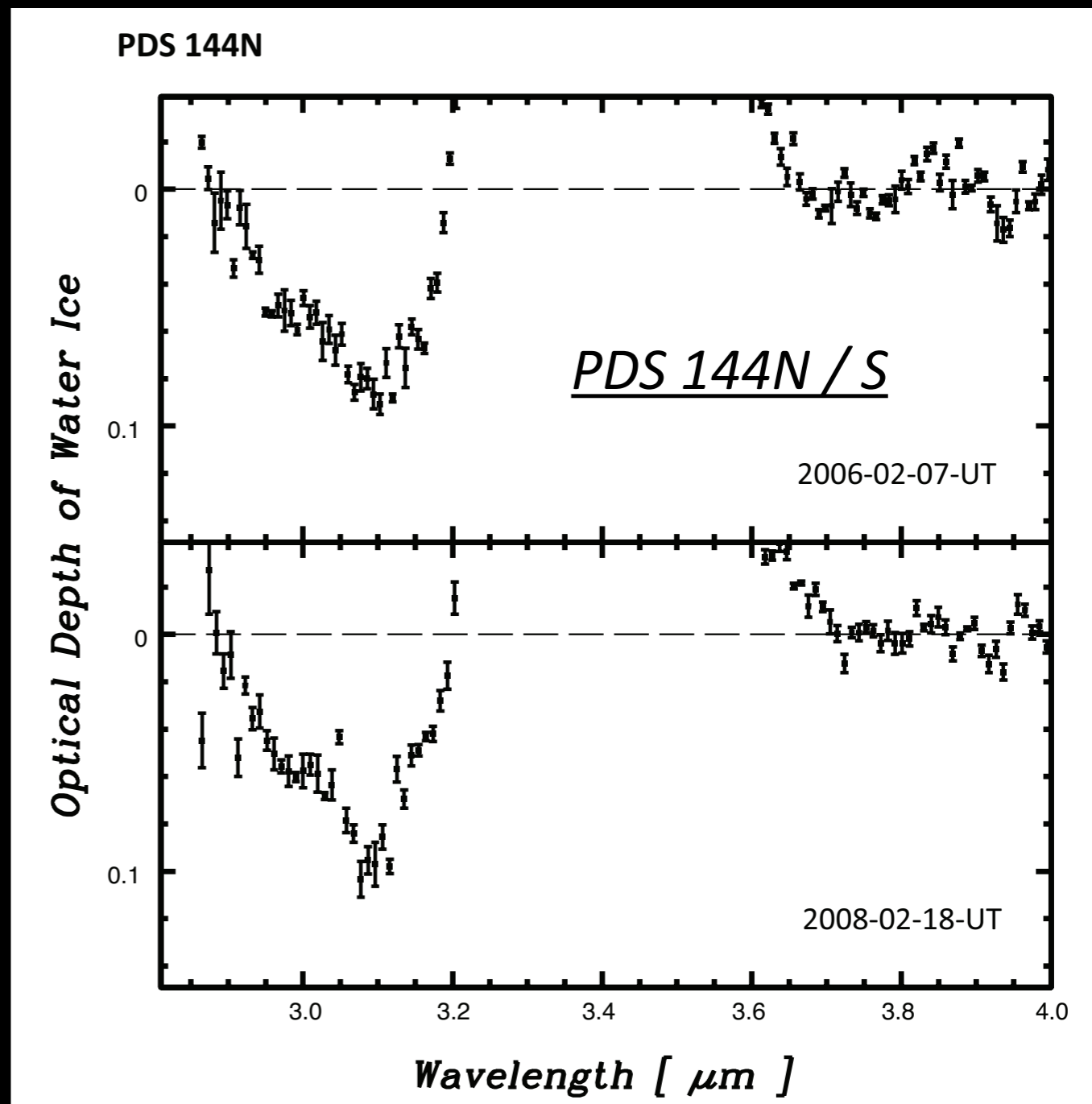
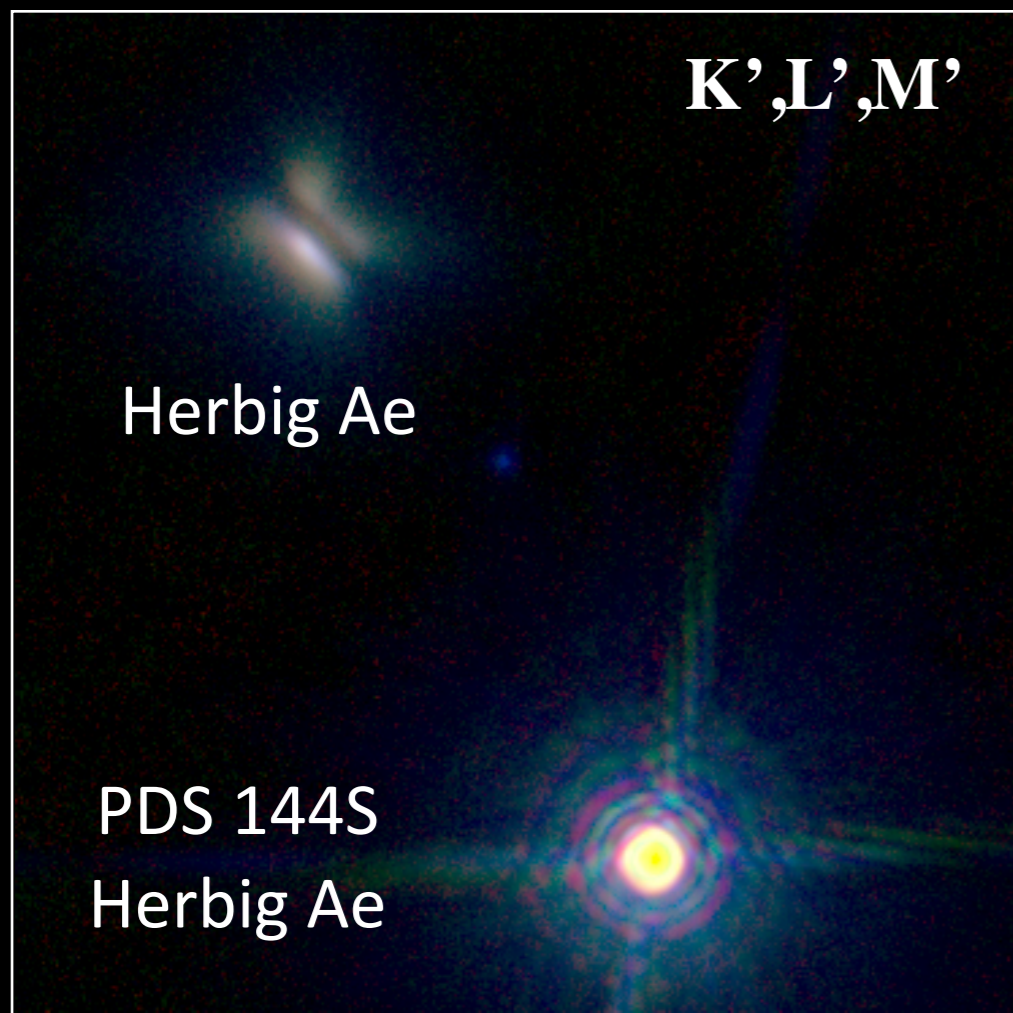
*gas*



# Benefit of Simultaneity

## Example of IRCS Result for "Ice"

PDS 144N

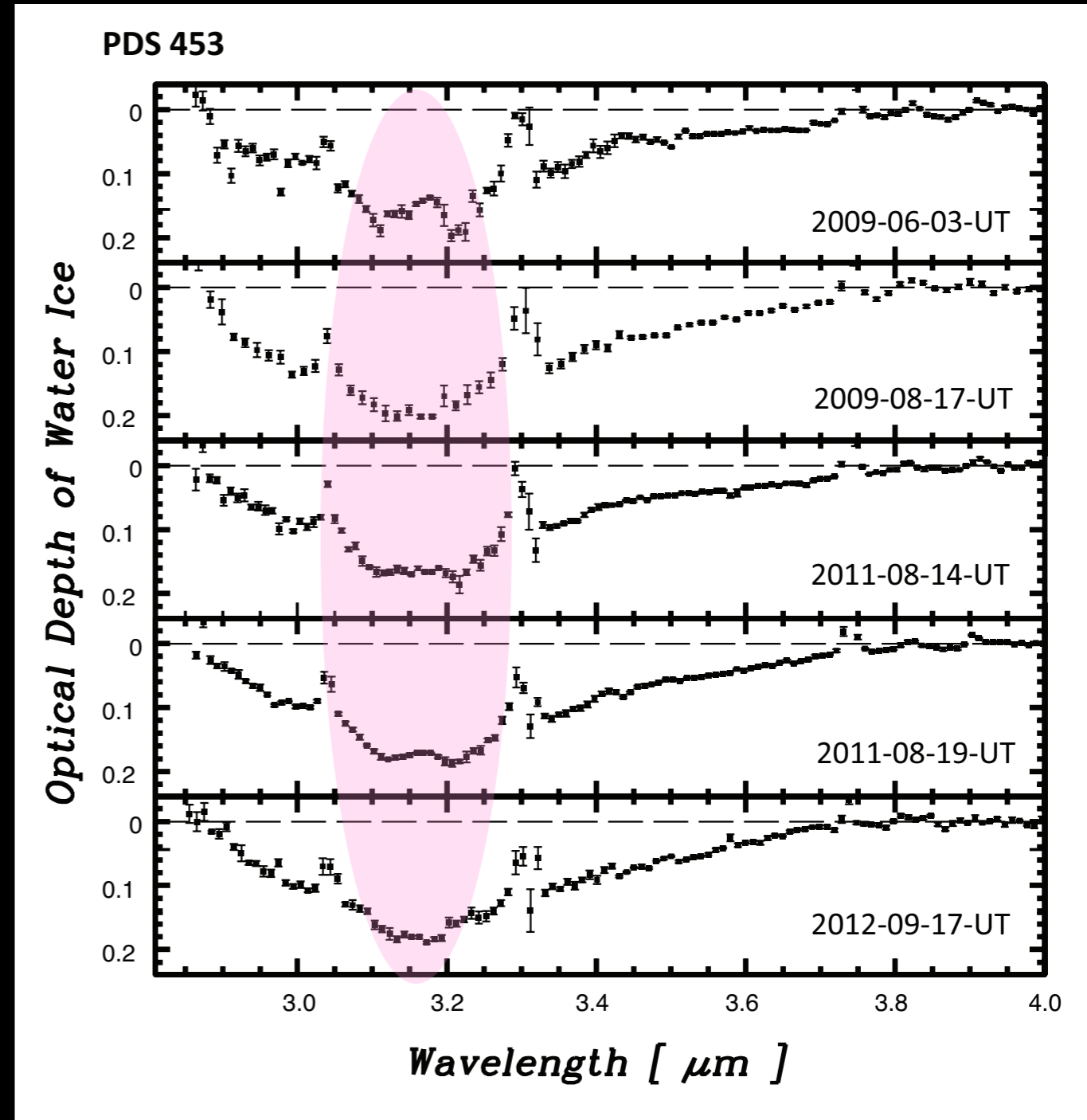
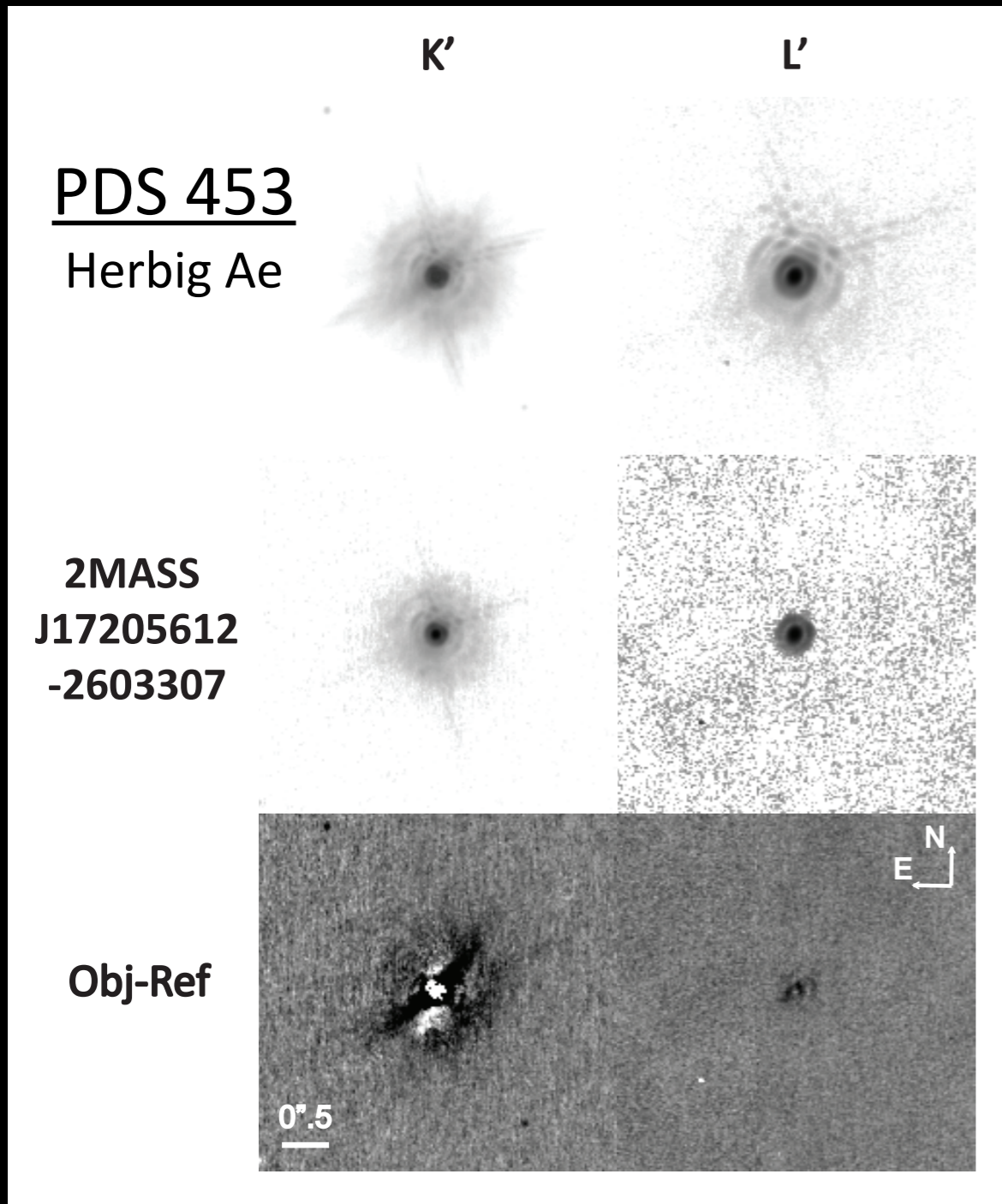


Terada+ in prep.



# Possible Target for "Ice" w/ MIMIZUKU Monitoring

IRCS+AO188 [NGS]



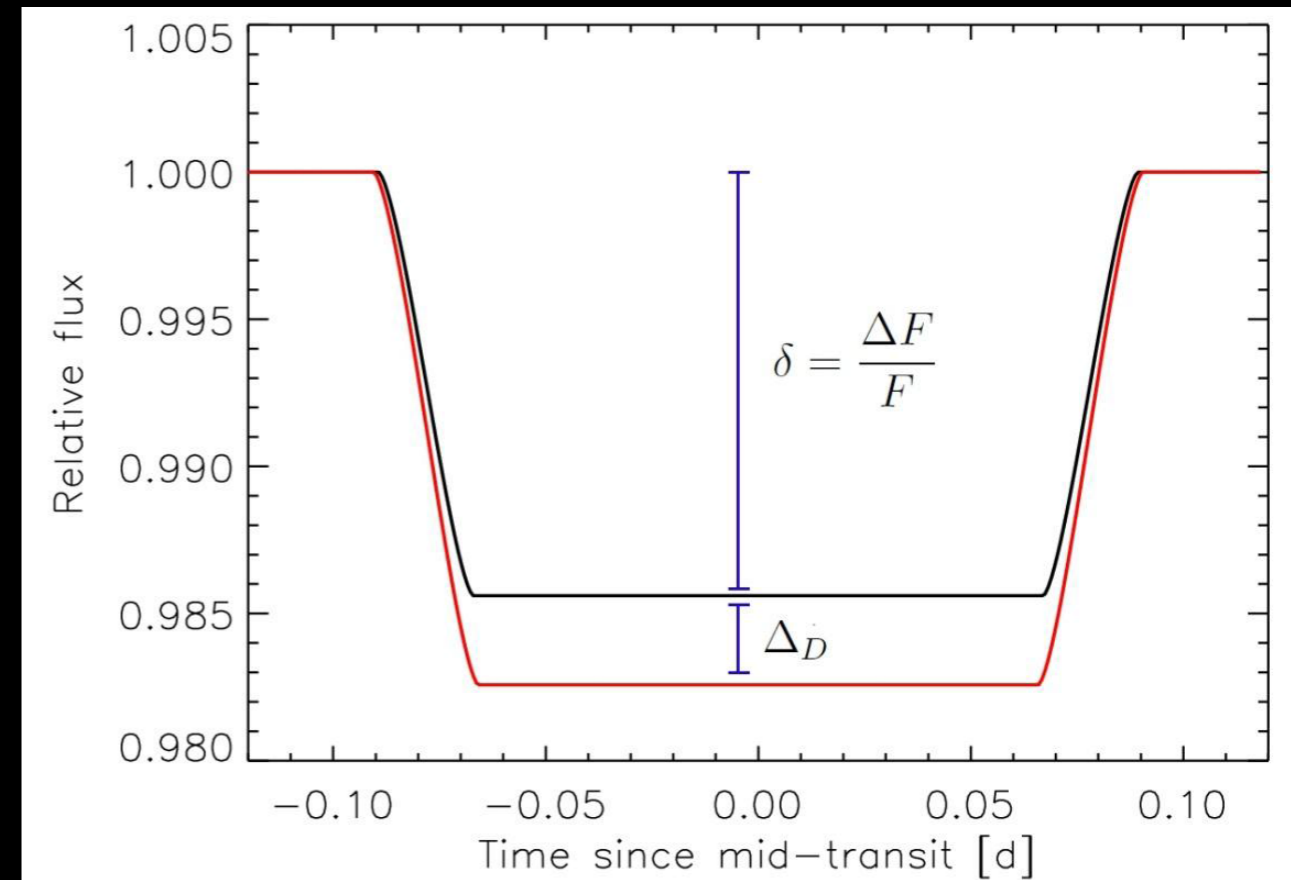
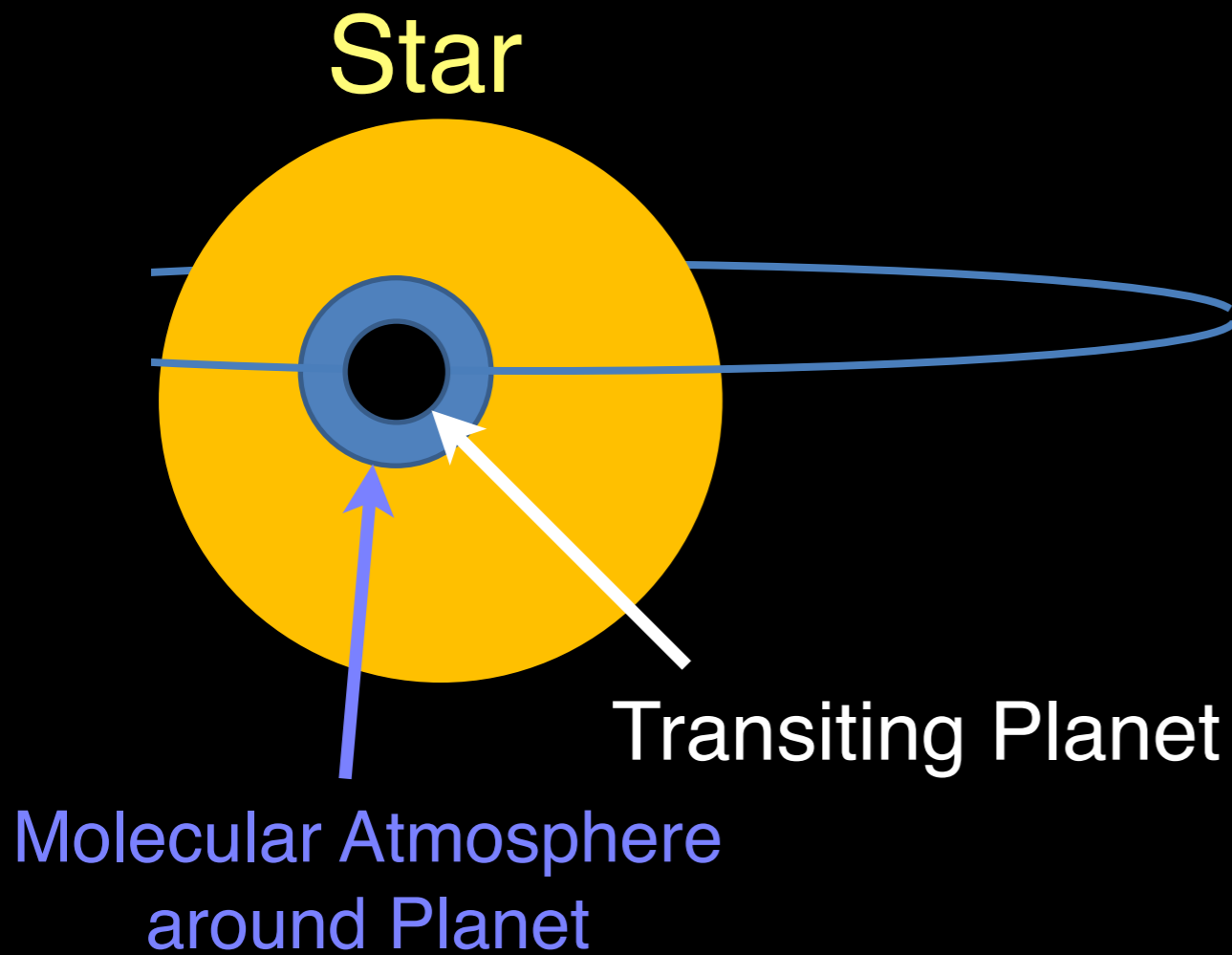
Terada+ in prep.

(II) Exo-Planetary  
Atmosphere in  
Transiting Systems



# Exo-Planetary Atmosphere in Transiting System

## *Theory for Detection*



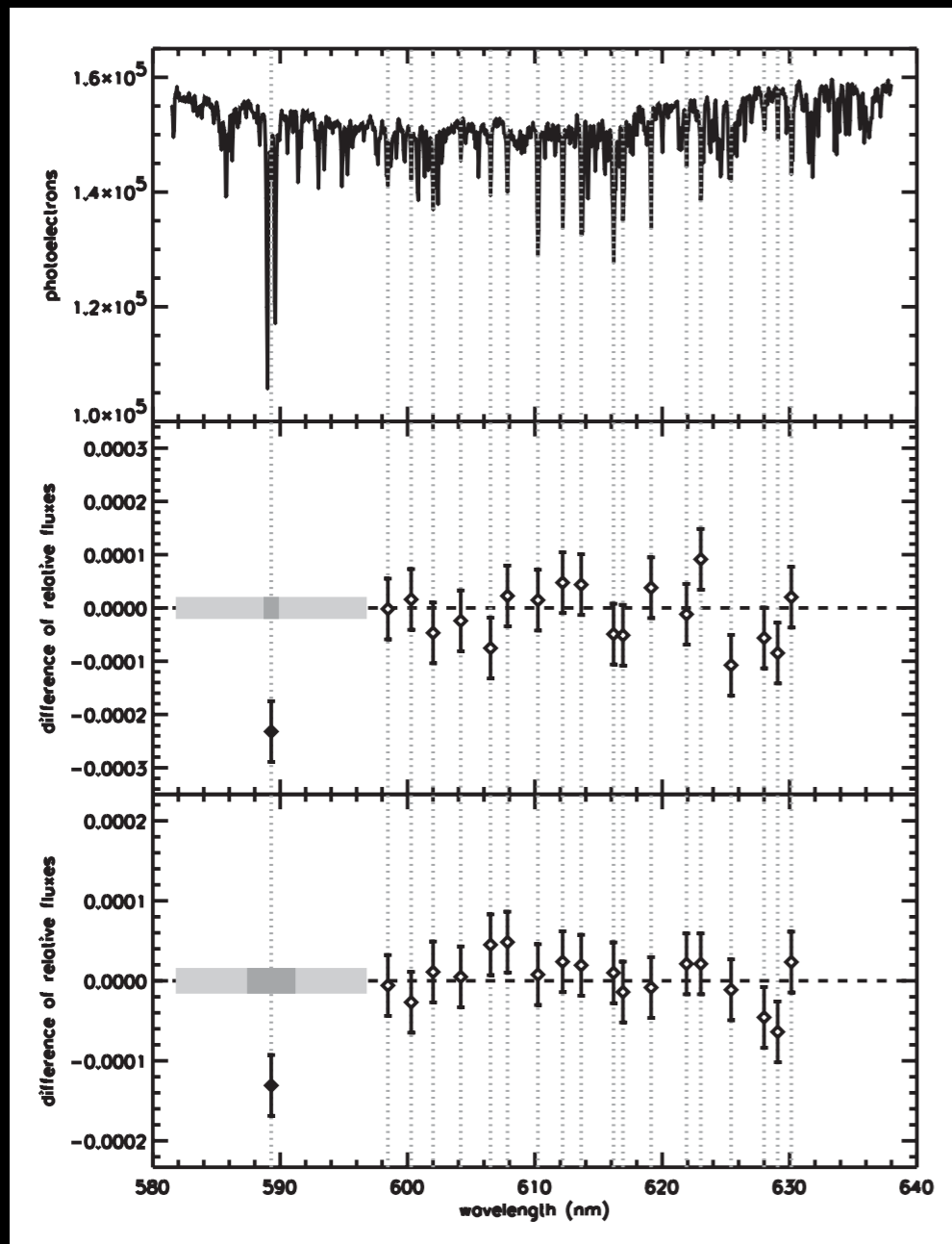
*Required Precision  $\sim 10^{-4}$*





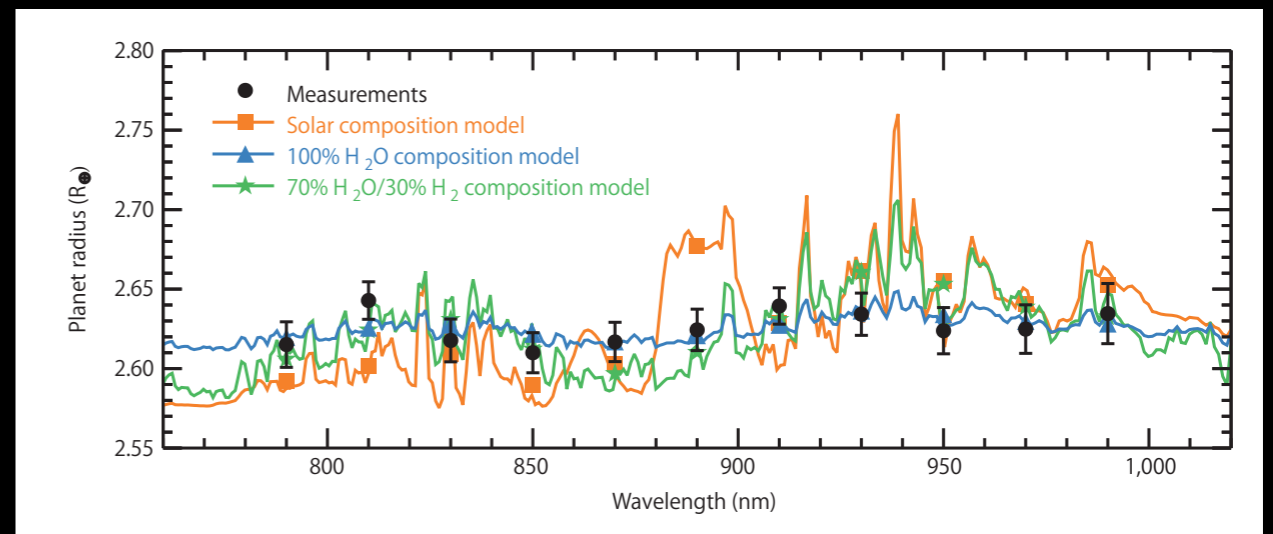
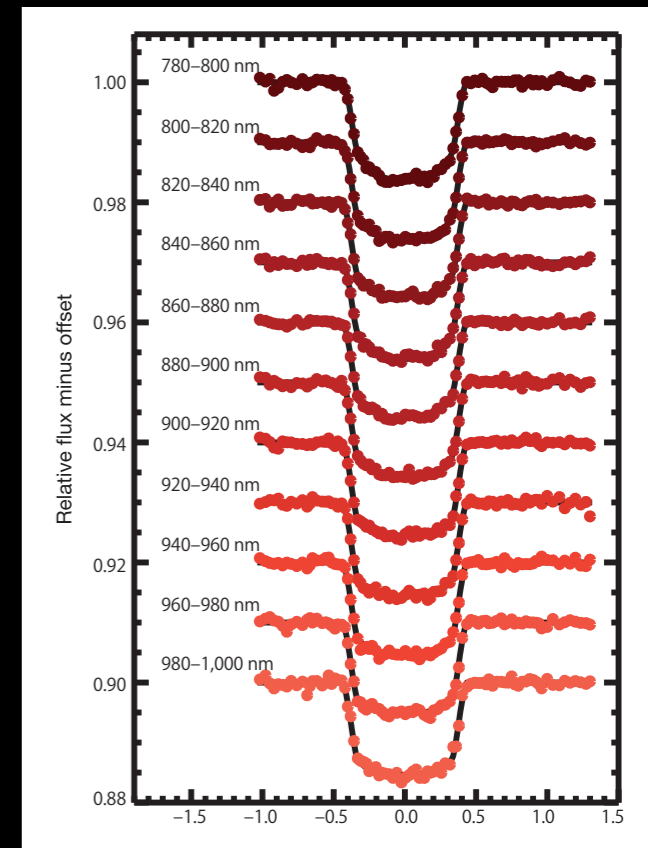
# Exo-Planetary Atmosphere in Transiting System: *Observational Achievements*

**First detection w/ HST STIS**  
NaD line: HD 209458



Charbonneau+ 2002

**Ground-based detection**  
w/ VLT FORS  
Optical Spectroscopy  
GJ 1214b

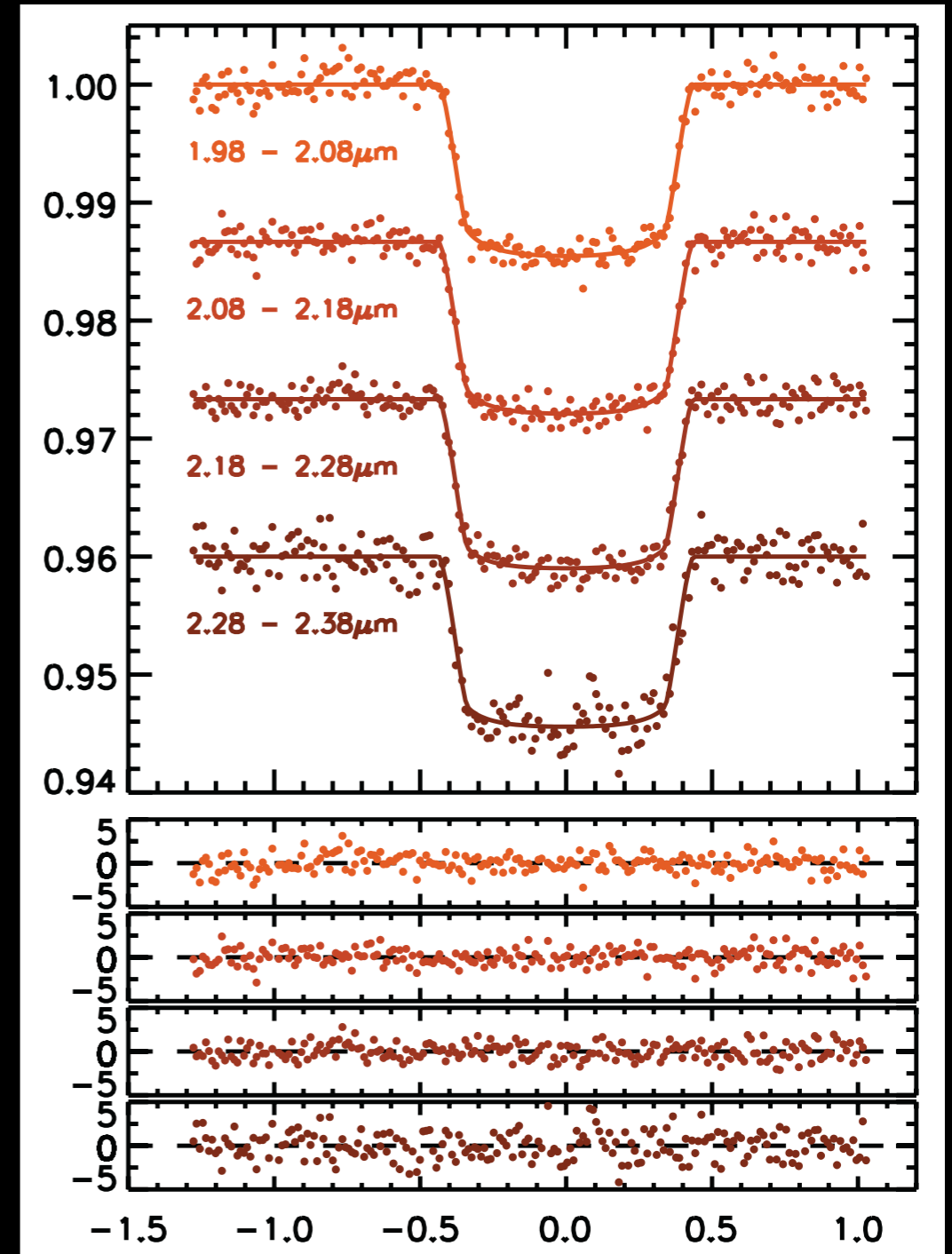
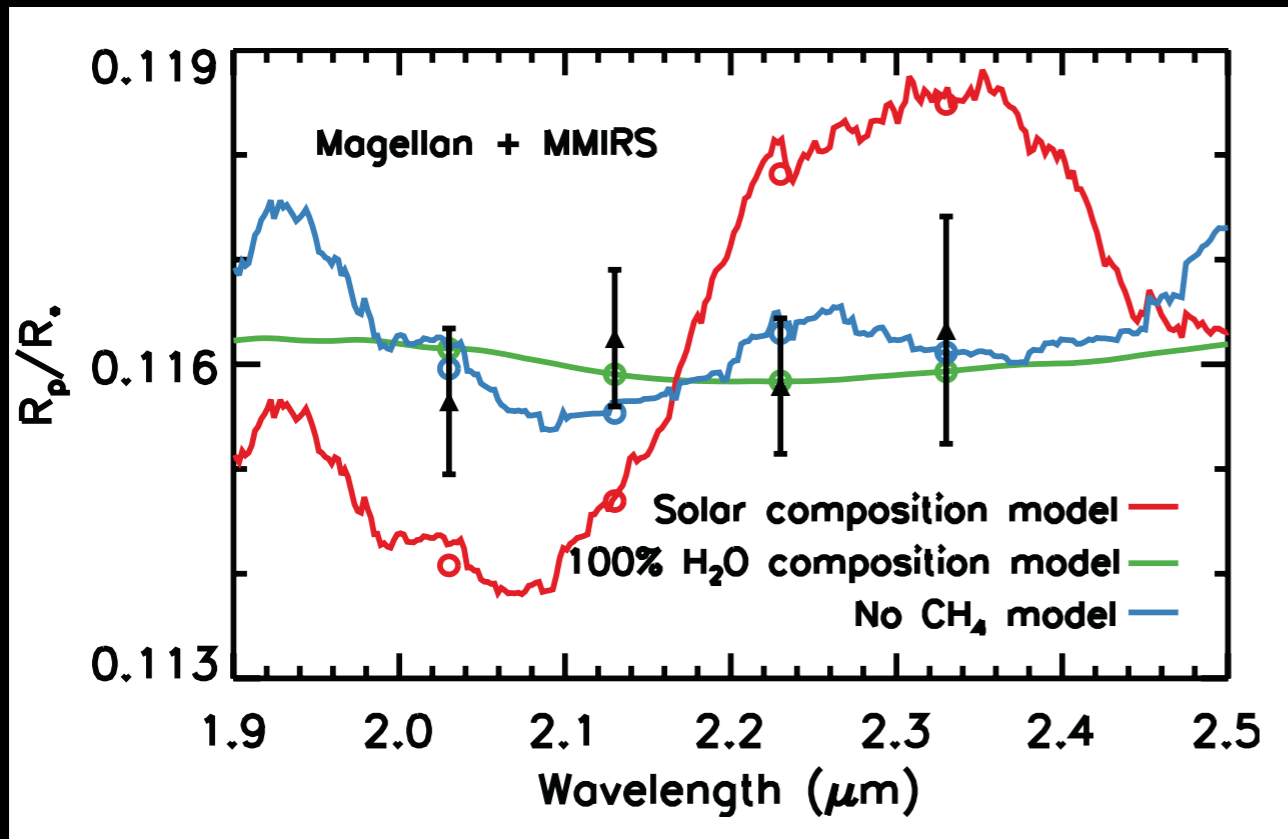


Bean+ 2010



# Exo-Planetary Atmosphere in Transiting System: *Ground-based Trial in IR*

**Ground-based IR detection  
w/ Magellan + MMIRS  
Infrared K-band Spectroscopy  
GJ 1214b**

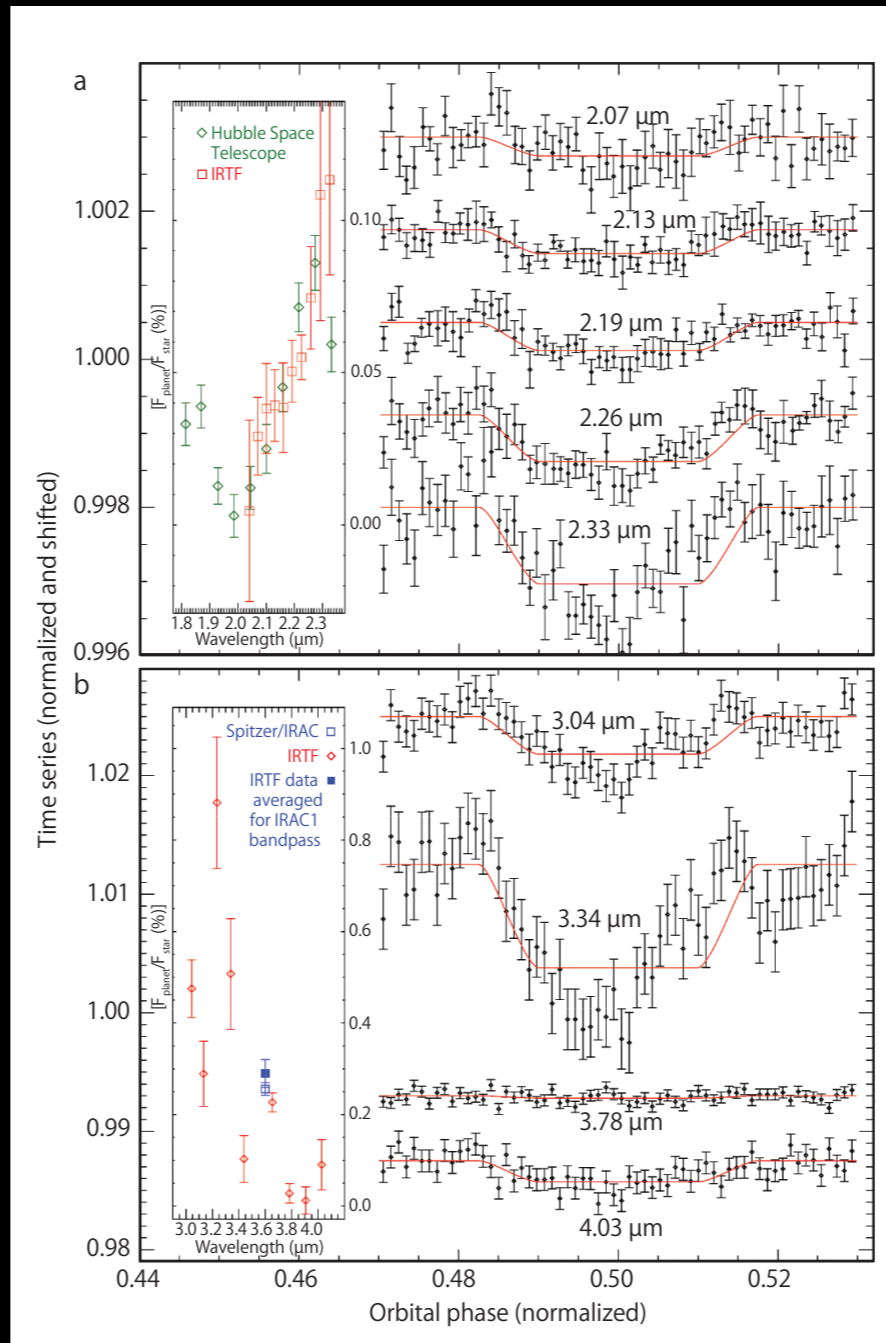
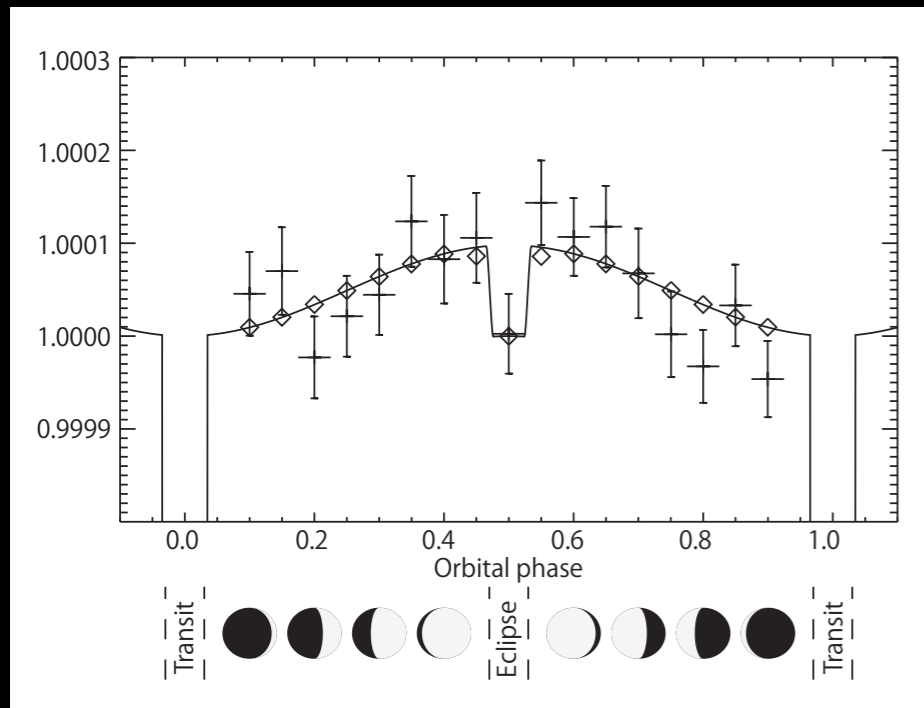


Bean+ 2011

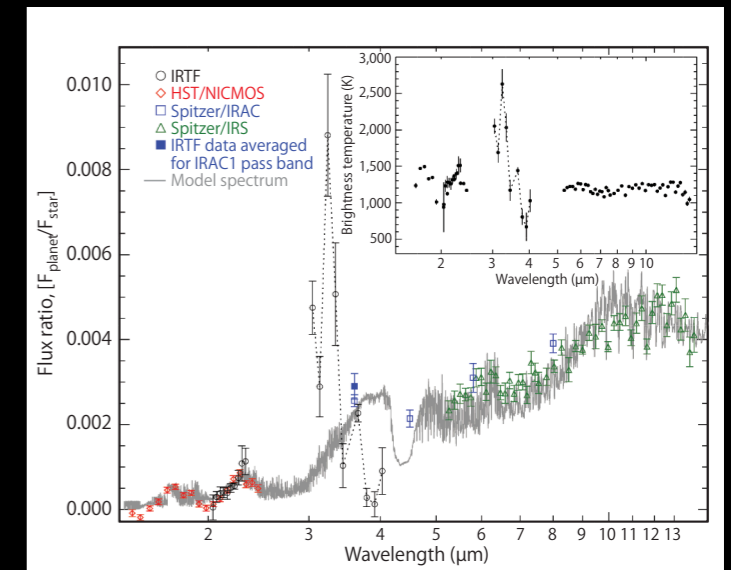


# Exo-Planetary Atmosphere in Transiting System: *Secondary Eclipse in 3μm*

## Secondary Eclipse



**Ground-based 3μm  
detection  
w/ IRTF + SpeX  
K,L-band Spectroscopy  
HD 189733**



Swain+ 2009



# Exo-Planetary Atmosphere in Transiting System:

*Issues: Keys to Achieve “10(-4)”*

**Especially for IR, it’s so difficult..**

- a. *Stick the target exactly at the same location of the detector.*  
=> *Uncertainty of “ununiform sensitivity” in the infrared detector has to be eliminated.*
- b. *Increase the S/N using over-sampling of the object.*  
=> *Usually defocused image is used.*
- c. *“Simultaneously” observe the reference.*  
=> *Eliminate variability of atmospheric absorption*



# Summary

- ✓ *Shorter wavelength region of L-band: one of the most fruitful area at the Atacama site.*
- ✓ *“Simultaneous spectroscopy” would be a key for detecting gas features and monitoring ice profiles in the protoplanetary disks.*
- ✓ *“Exo-Planetary Atmosphere” in the secondary eclipse may be able to be detected in  $3\mu\text{m}$ . Stability of guiding should be an issue.*
- ✓ *MIMIZUKU would work for “simultaneous spectroscopy” at Subaru complementarily with RAVEN (+IRCS).*

*# RAVEN 2nd Science Meeting will be held on*

*# July 25 and 26, 2013 at Waikoloa, Hawaii,*

*# which is right after TMT forum. Your interest is welcome!*

