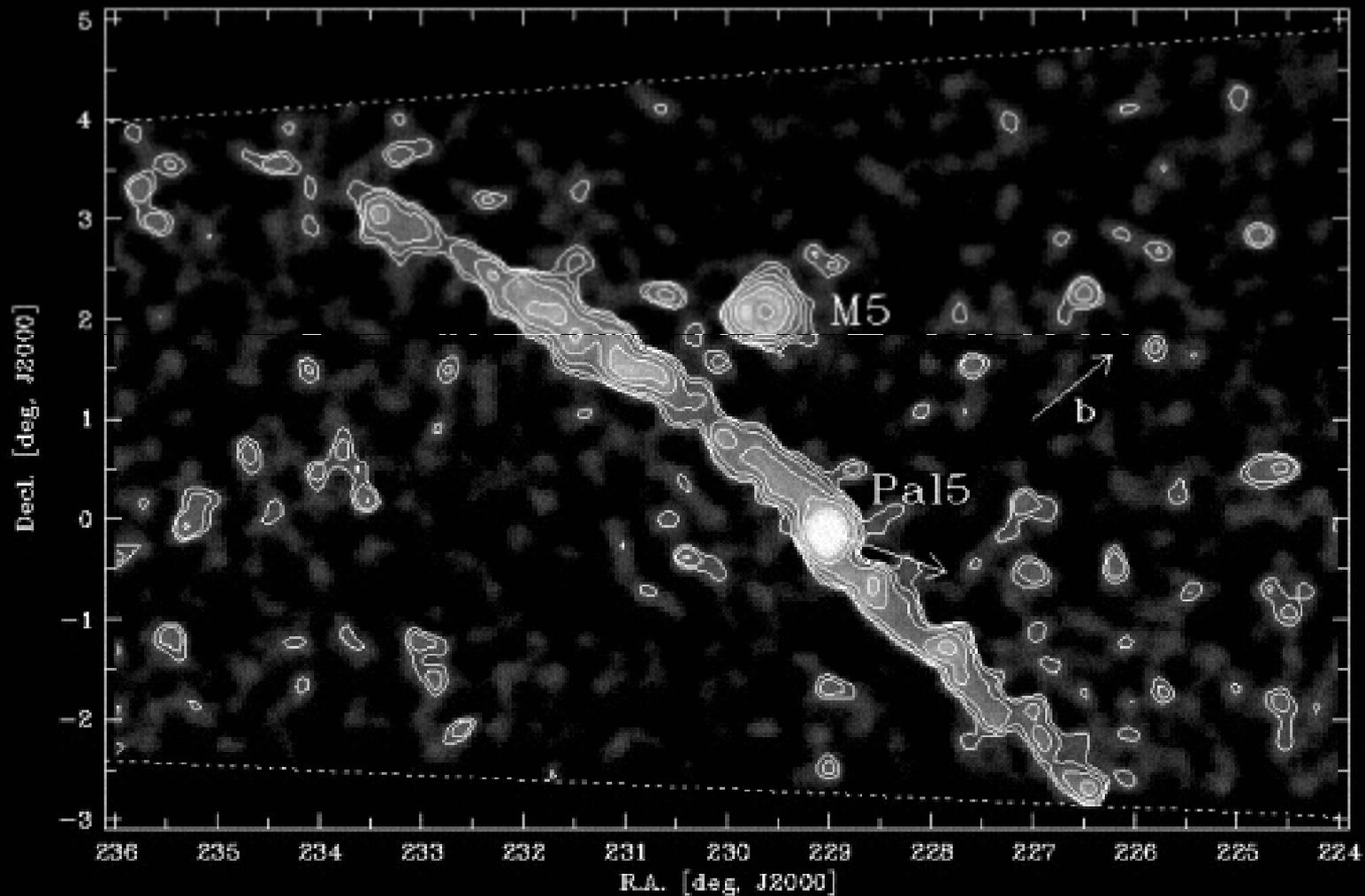


# 銀河考古学 I 銀河系



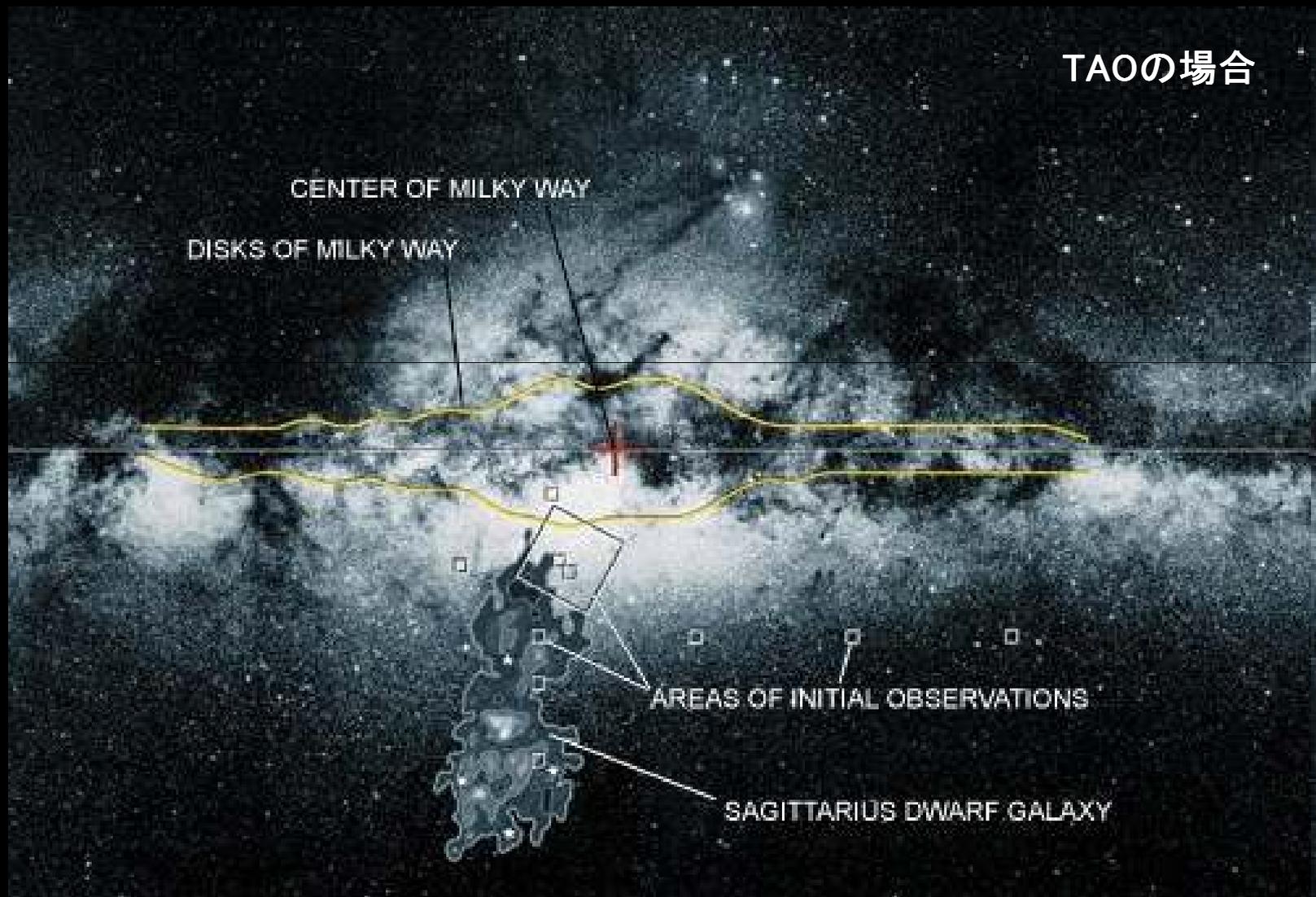
**N.Arimoto**  
**NAOJ (Japan)**

# 近红外多体分光装置

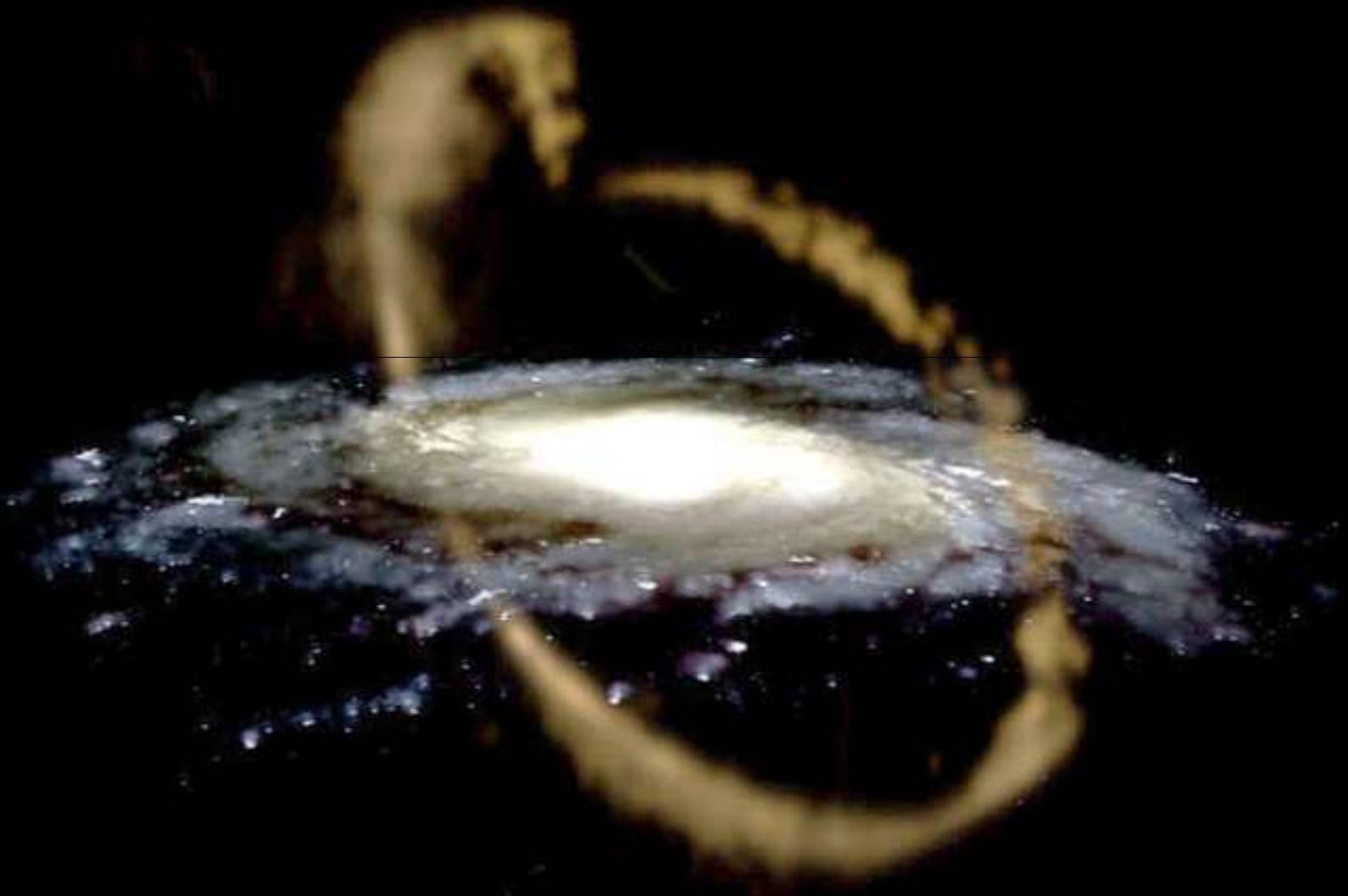


# いて座矮小銀河

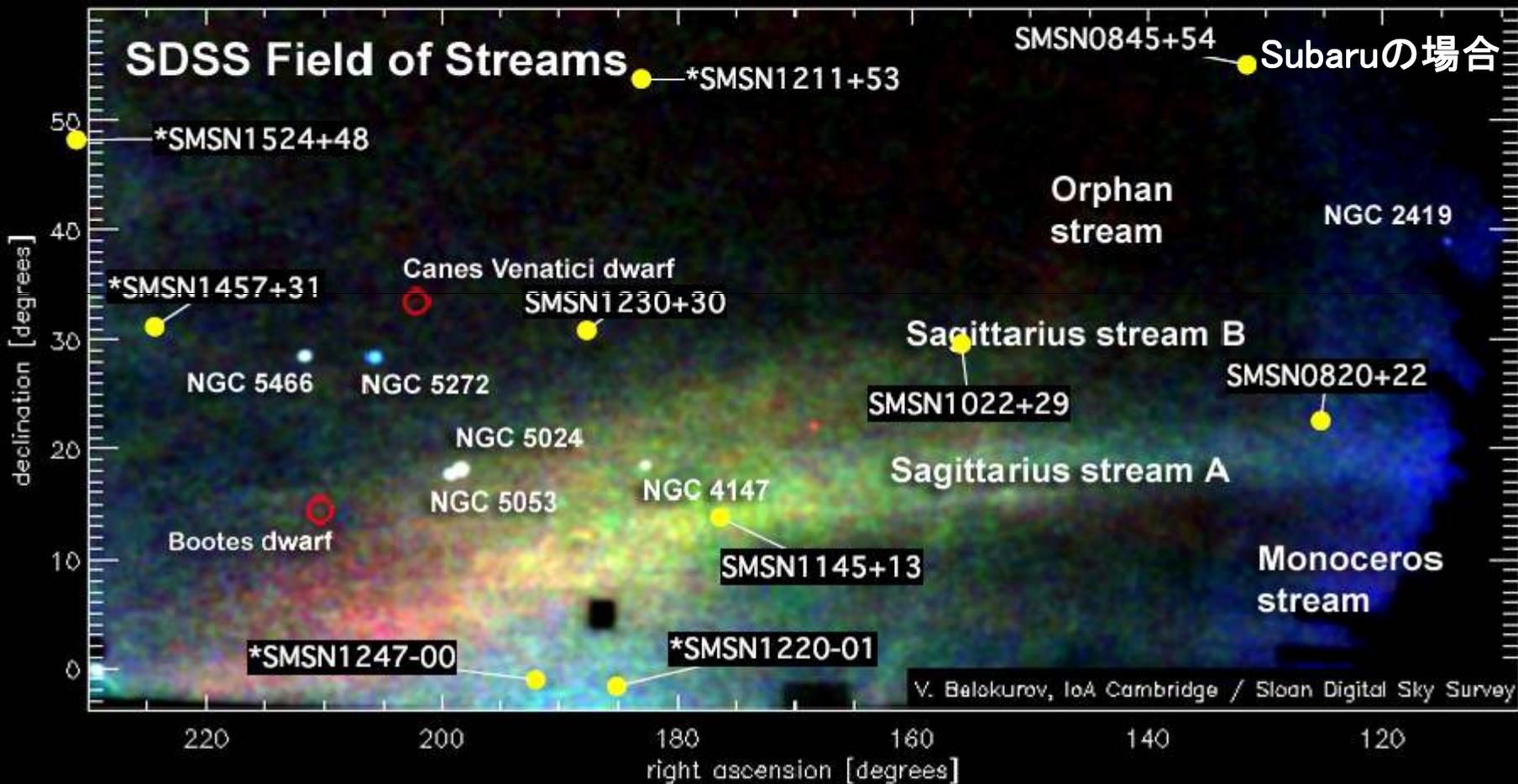
TAOの場合



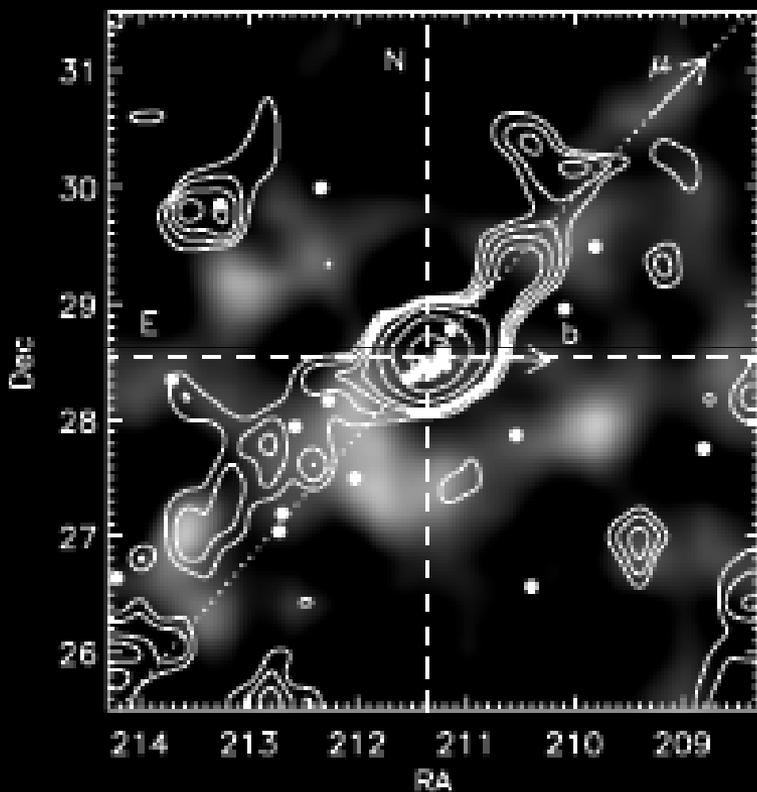
# いて座矮小銀河



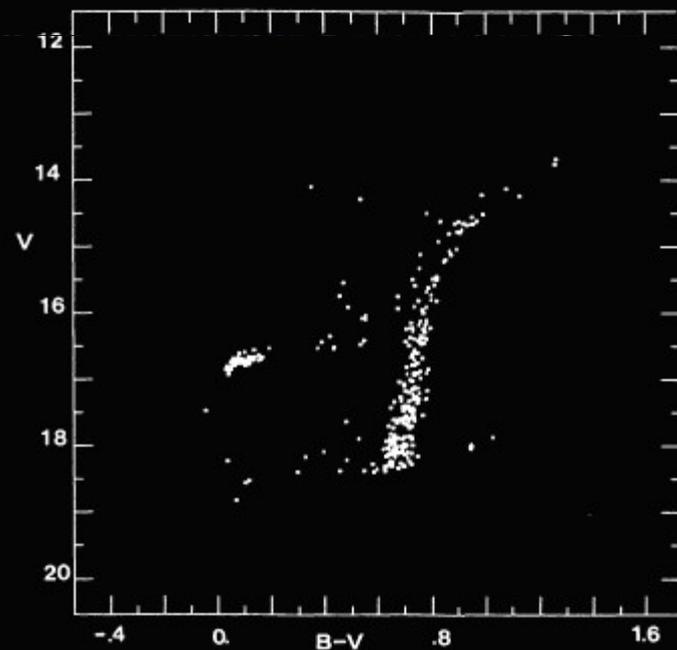
# 銀河系のストリーム構造



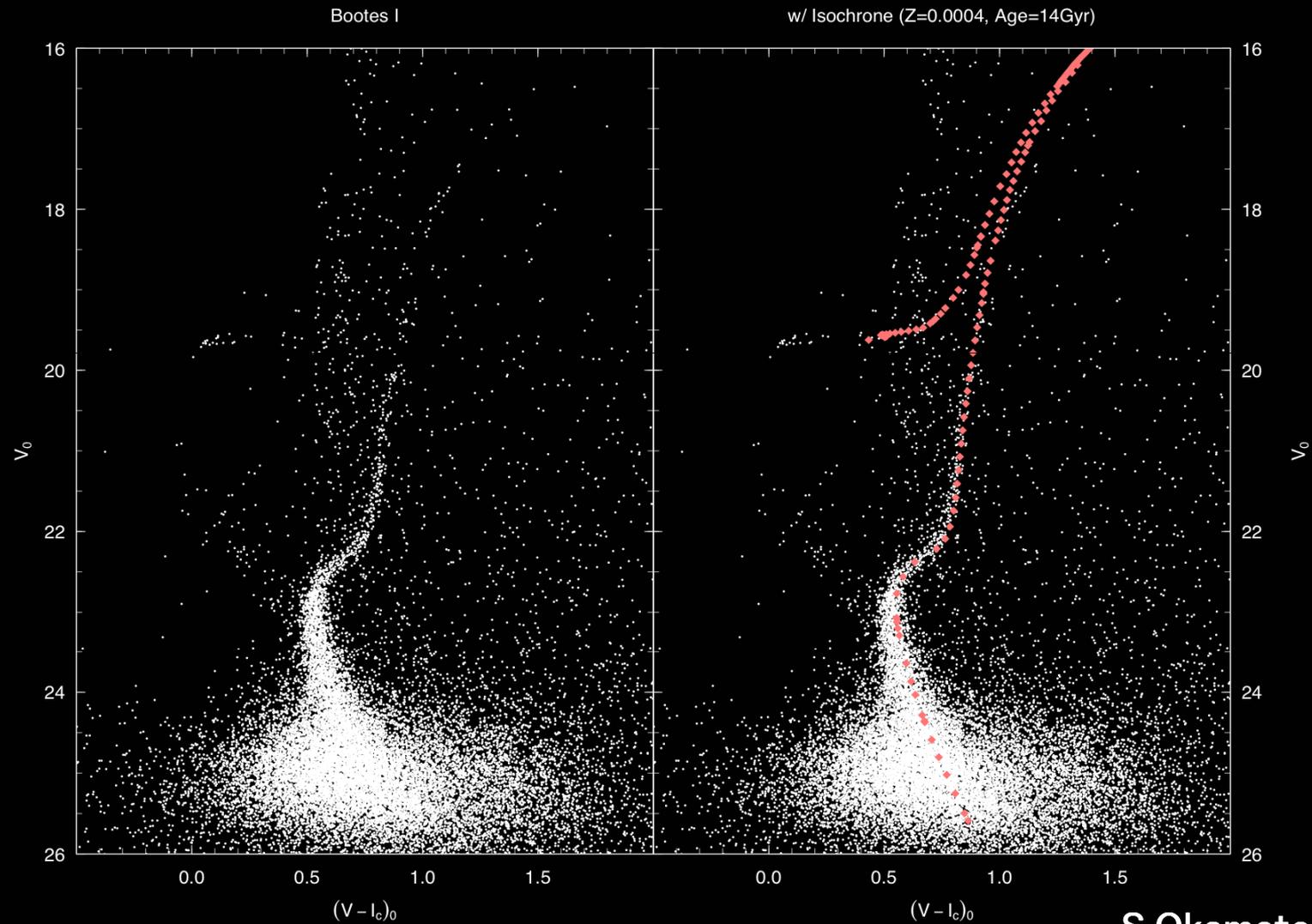
# ストリーム近傍の球状星団



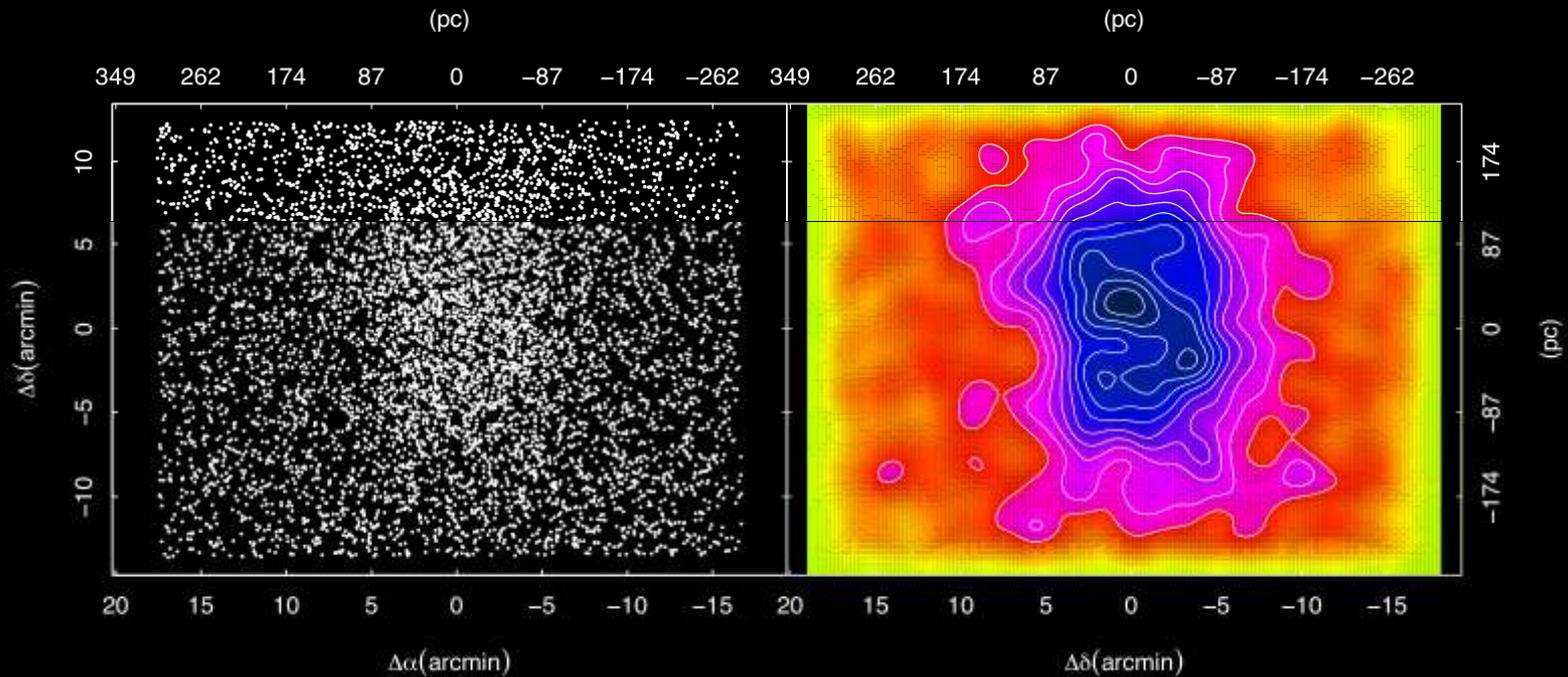
NGC5466 (Fellhauer et al. 2007)



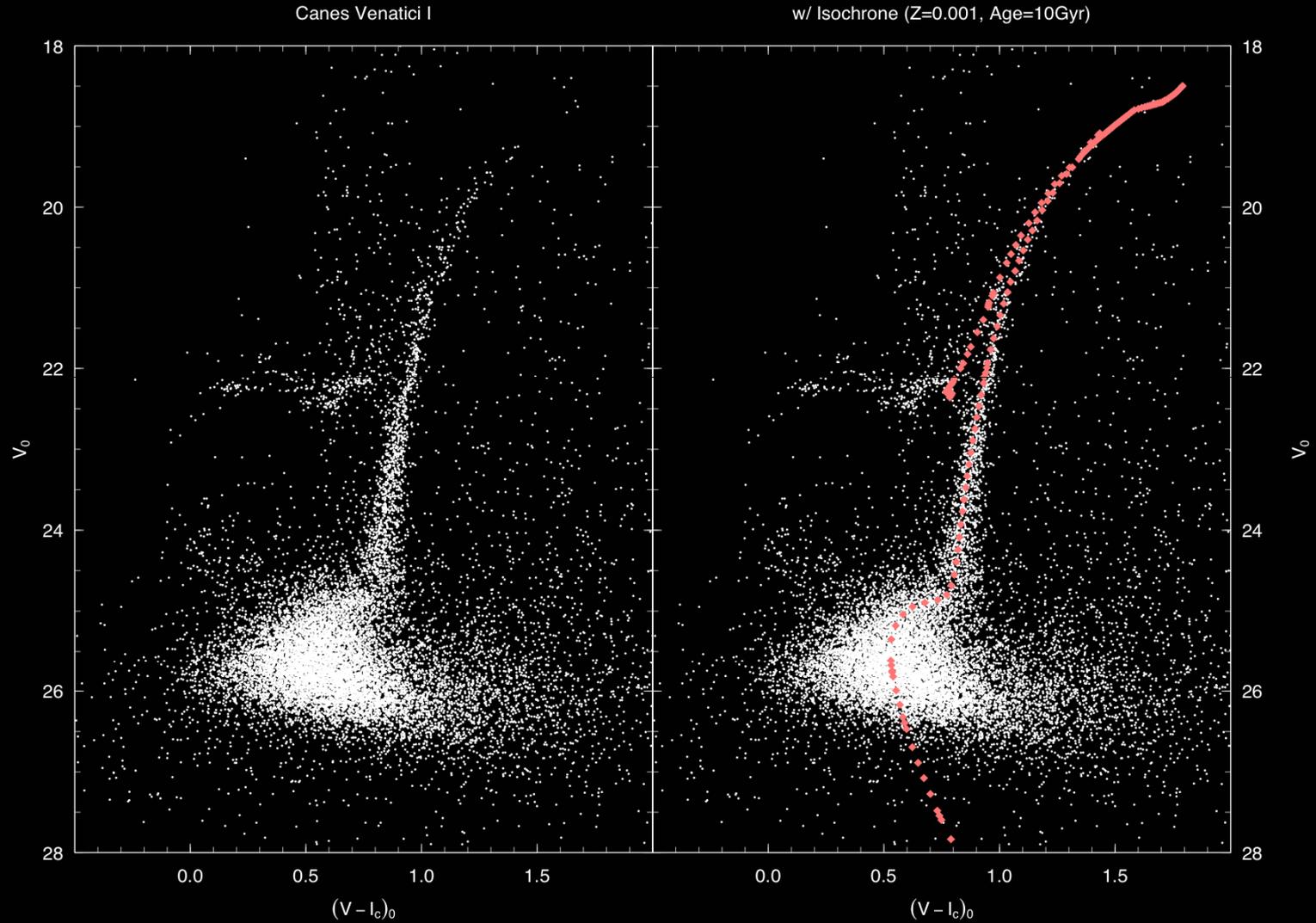
# Bootes I (UFdG)



# Bootes Iの空間構造

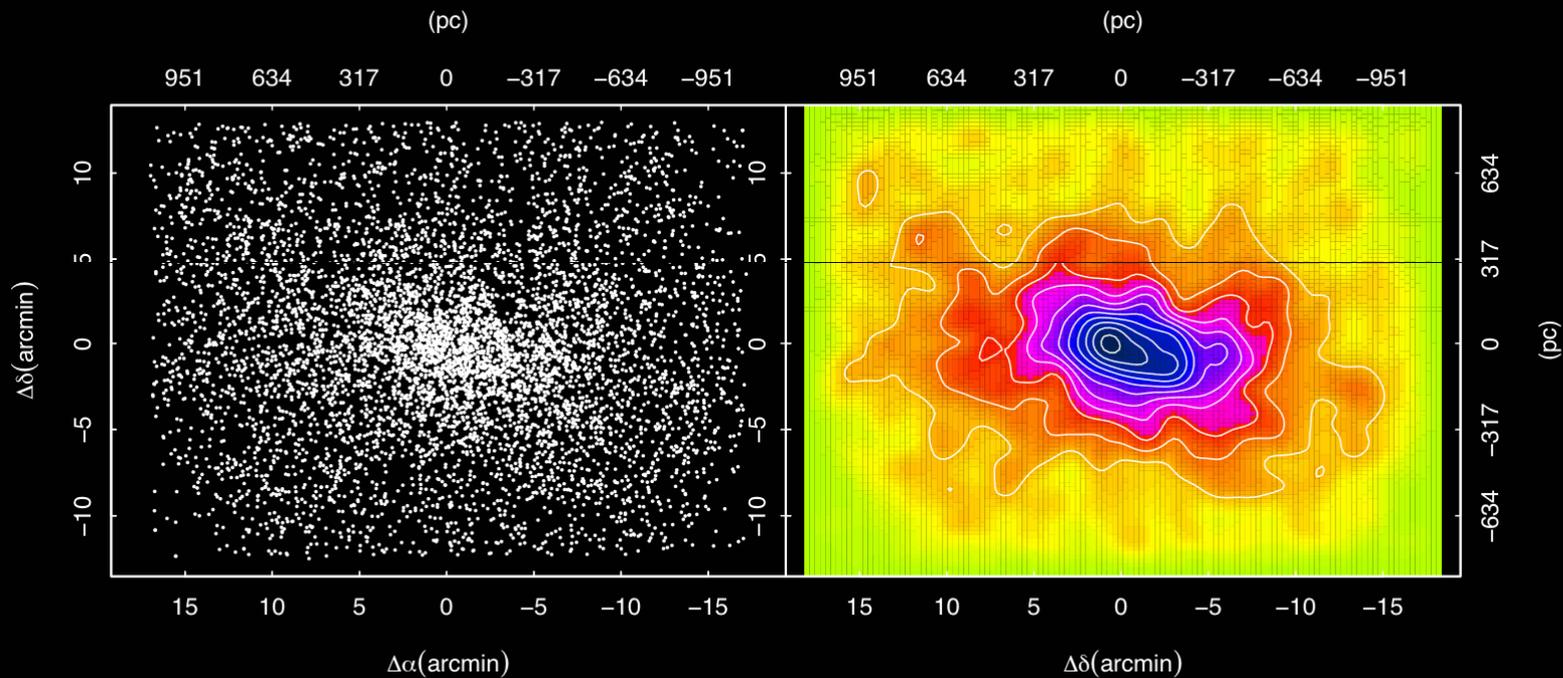


# Canes Venatici I (UFdG)

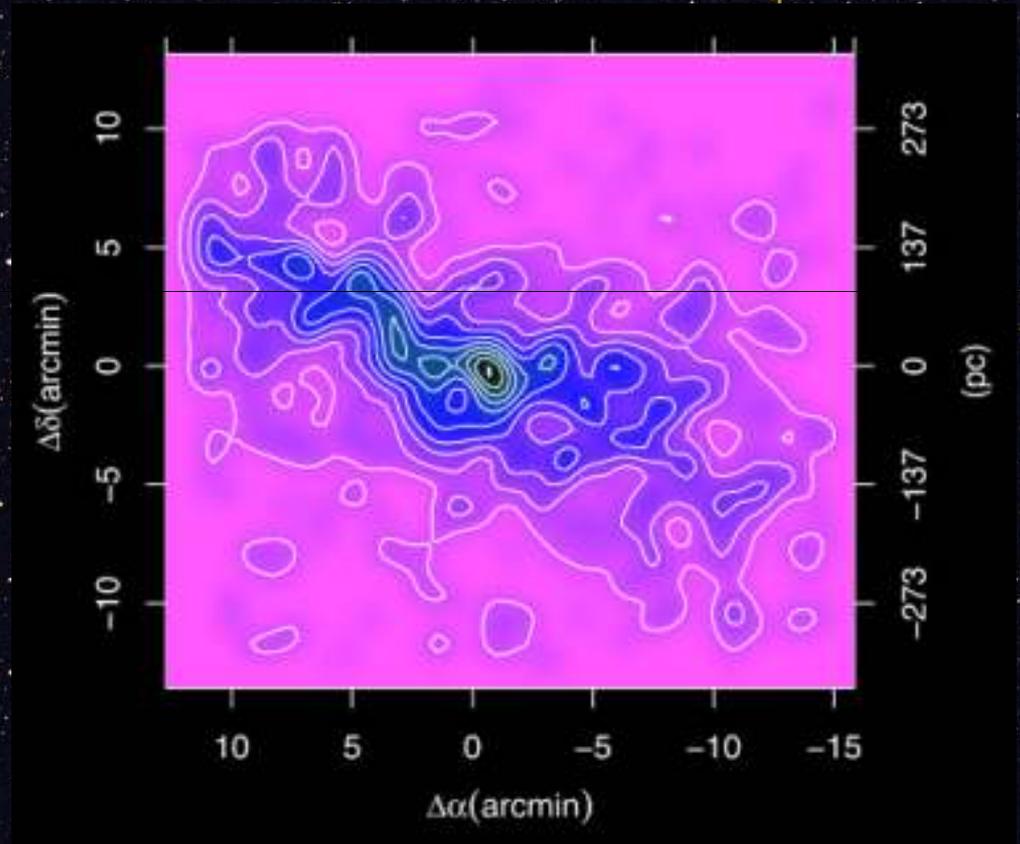
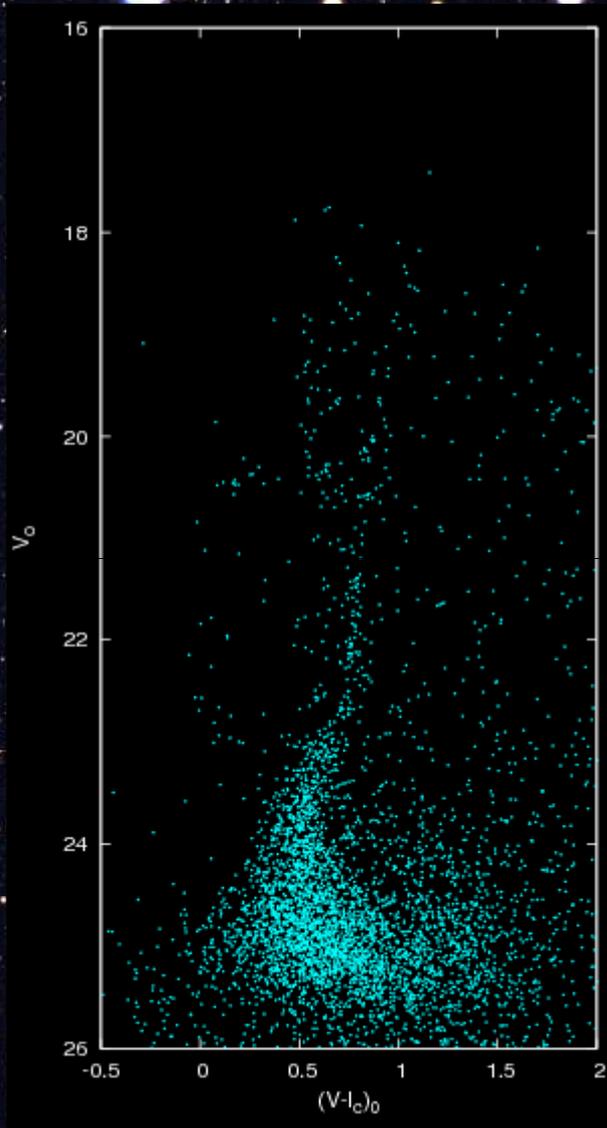


S.Okamoto (2009)

# Canes Venatici Iの空間構造

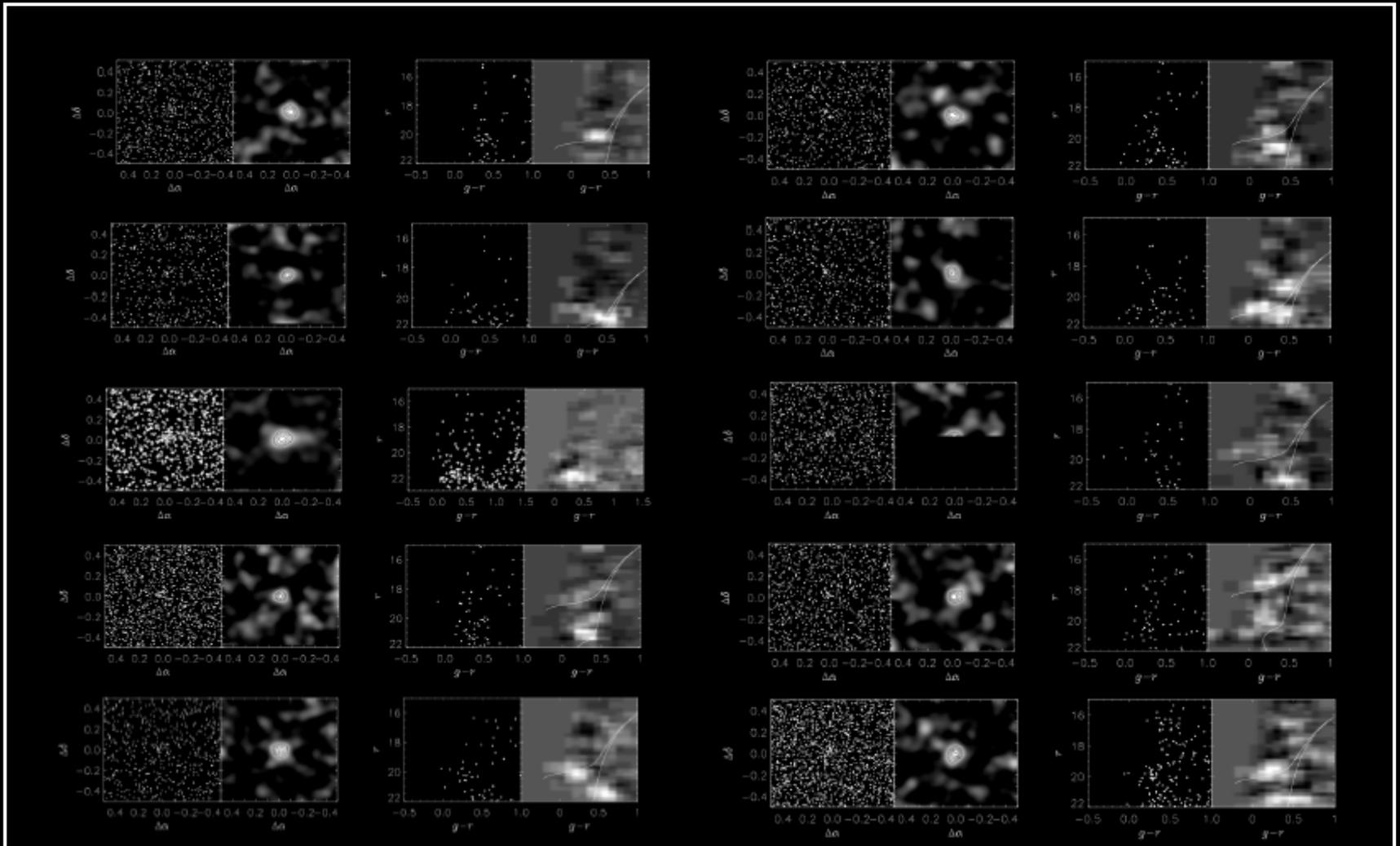


Okamoto et al. (2008)



Ursa Major (Suprime-Cam)

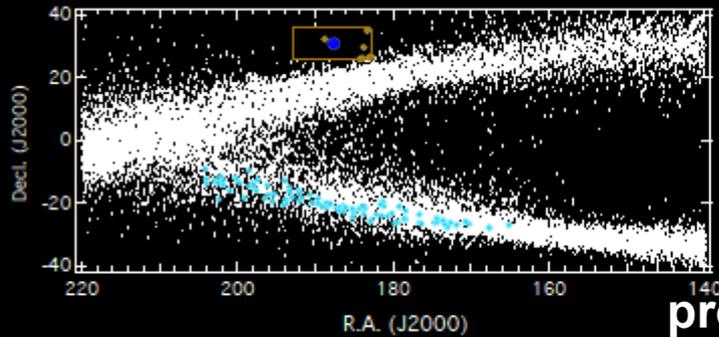
# Search for New UFGs



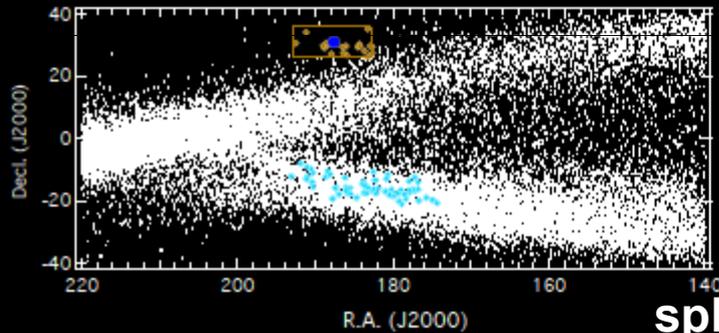
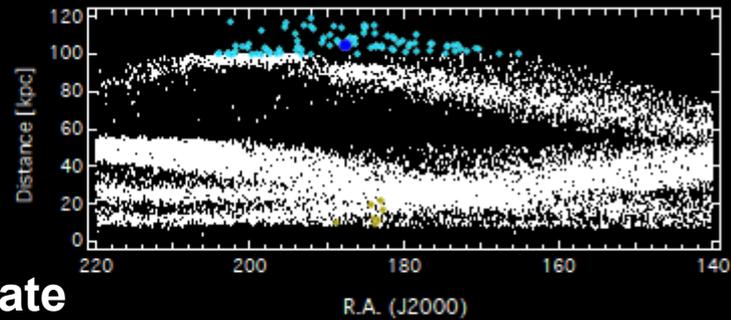
SDSS Candidates

ぎんがモデルとのひかく

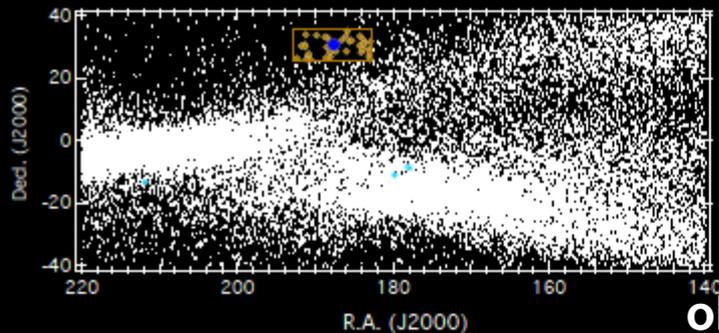
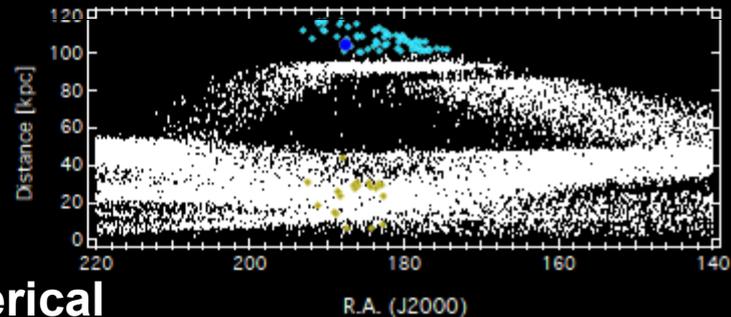
# ストリームの視線速度・距離



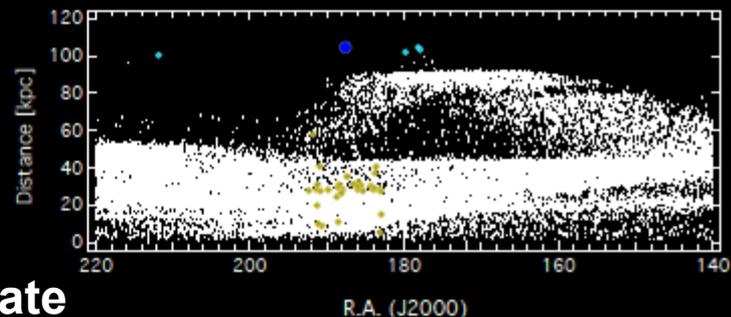
prolate



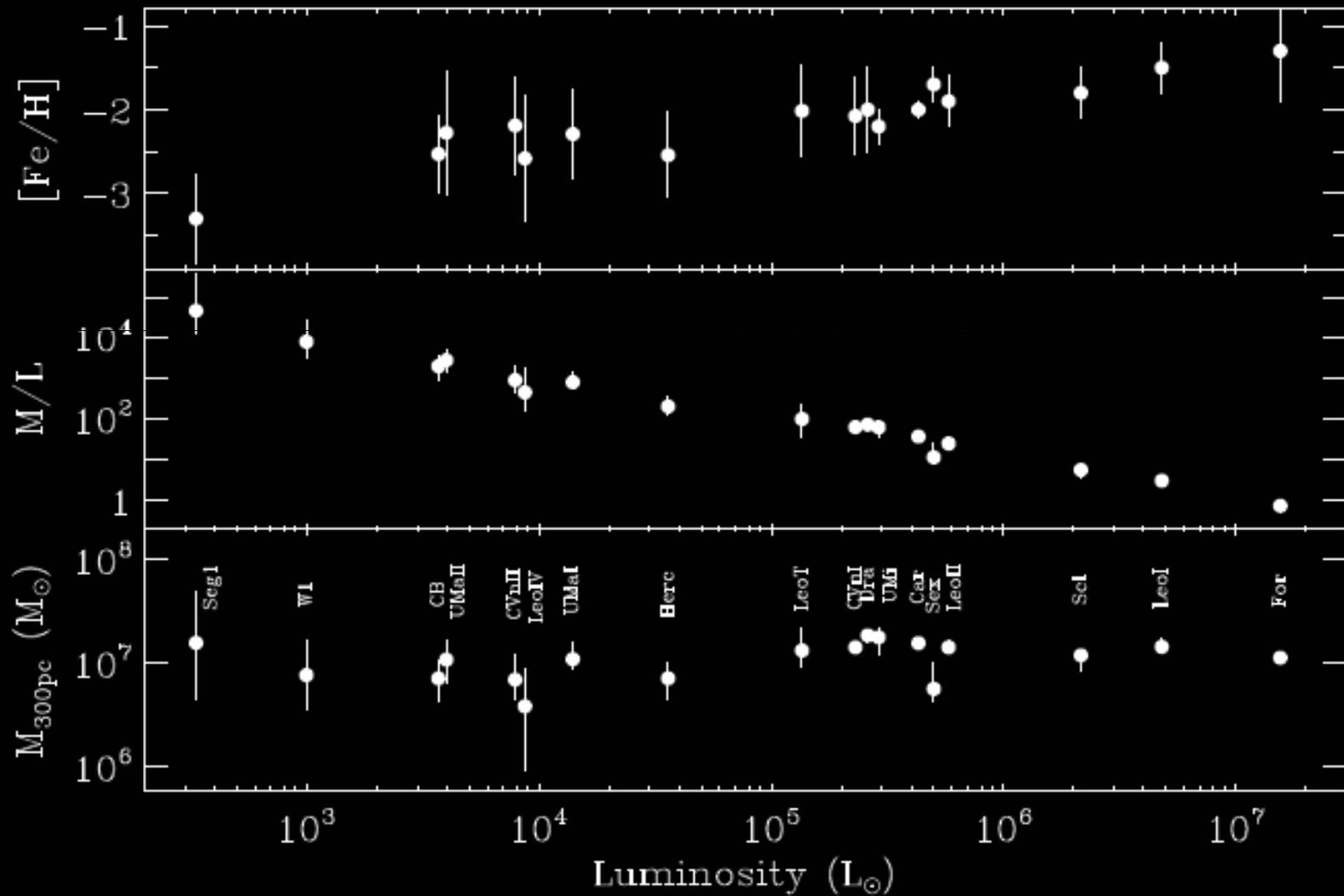
spherical



oblate

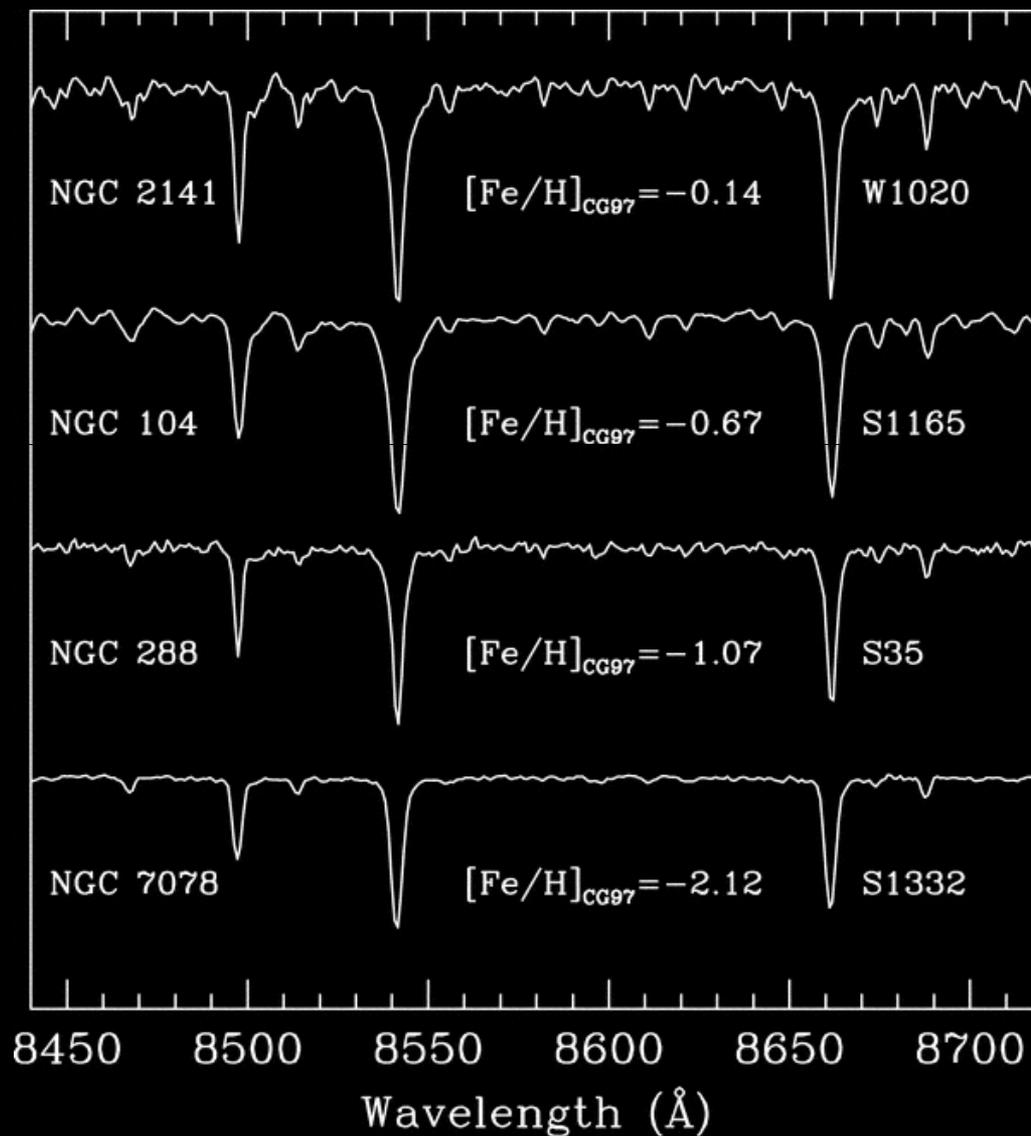


# ストリーム星の金属量

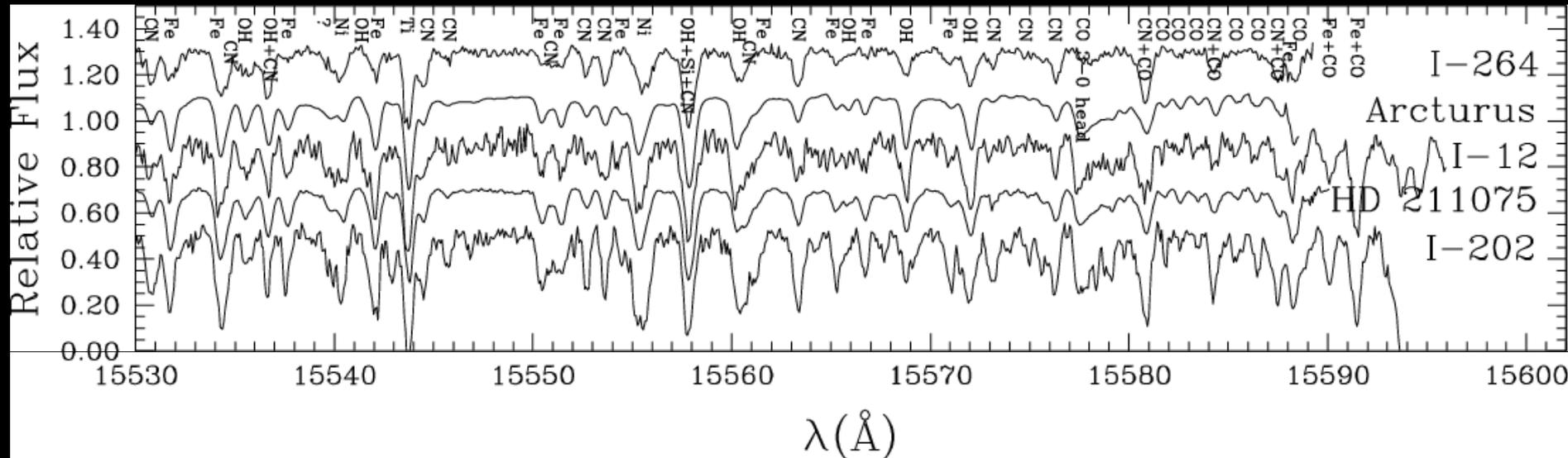


Call Triplet as Metallicity Indicator

# 近赤外での金属吸収線



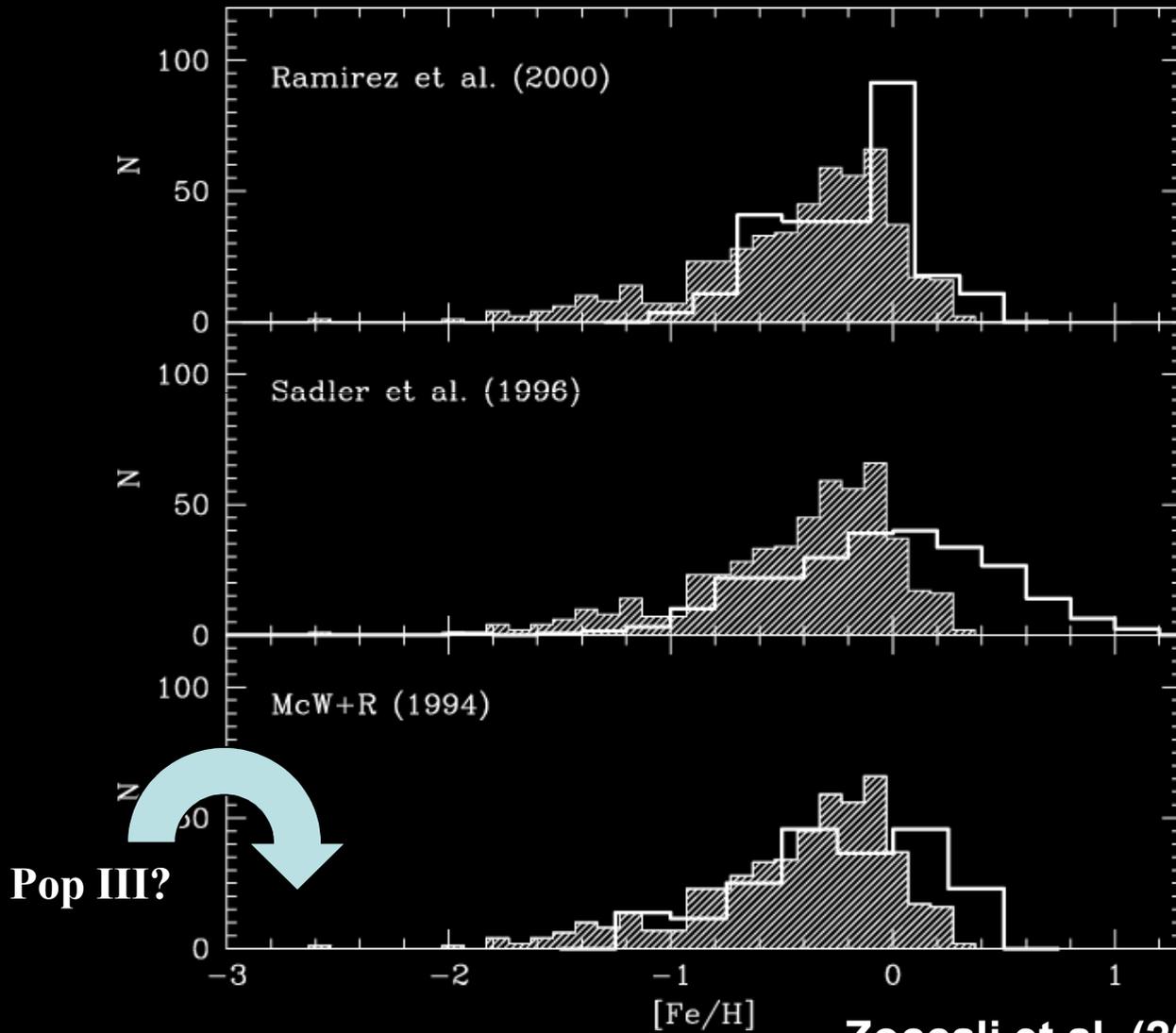
# 近赤外吸収線



Gemini-S/Phoenix (R=50000)

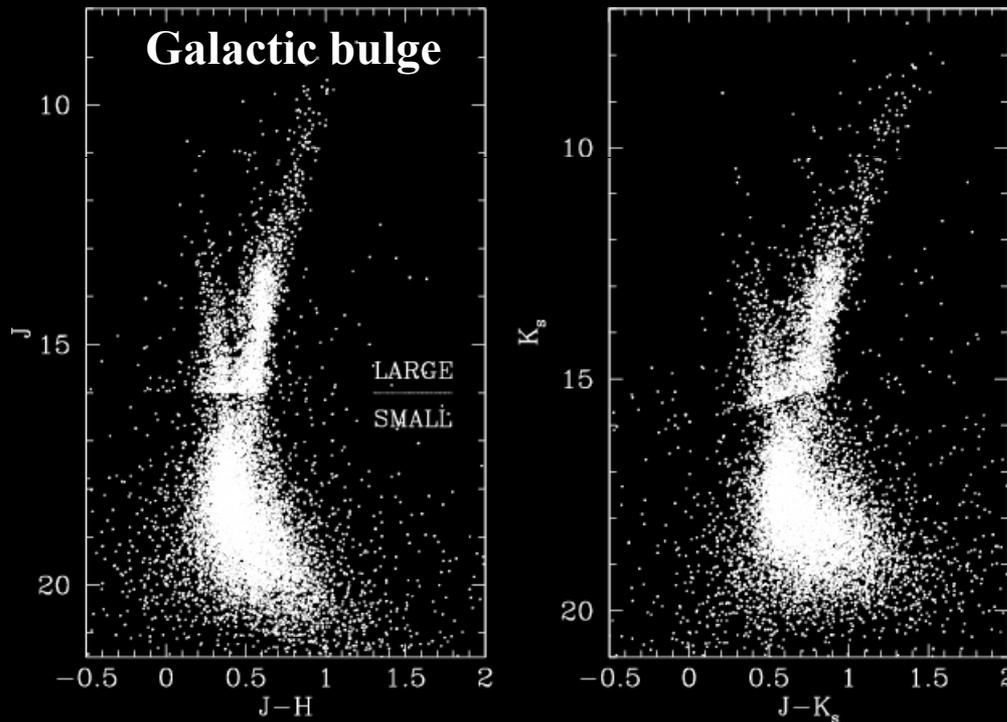
- 近赤外には原子やイオンの吸収線があまりない。
- あっても高電離の金属の吸収線で複雑で解釈も難しい。
- 多くの吸収線は正しく同定されておらず、吸収線の強度も不明である。
- その代わりに分子の吸収線があるが、TiOやZrOの吸収線はない。
- 分光器の性能は可視光域に比べて劣っている。
- 近赤外データだけから星のパラメータを決定するのは困難である。

# Population III stars in Galactic Bulge



# Specific Goals of HR Survey

- **Formation histories of galaxies:** stellar population study in terms of colour-magnitude diagrams (CMDs) and abundance patterns can reveal how bulge, disc and dSph's formed and evolved.



## Targeted Regions

bulges, disc, outer disc,  
halo, large area dSphs  
disrupted globular clusters  
M31/M33 large area  
(substructure in haloes)  
Galactic streams  
dwarf irregulars

# 銀河考古学 II 局所銀河群

